UNIVERSIDAD AUTONOMA
DE MADRID

Gravitational Waves
from the
EW Phase Transition

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14/11/19

The EW Phase Transition

WHY?

- ➔ Yield Precise Understanding of EWSB in Early Universe

The EW Phase Transition

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➔ Yield Precise Understanding of EWSB in Early Universe

➔ **(Possible) Answer to Open Mysteries at Interface of Particle Physics & Cosmology**

Origin of Matter-Antimatter Asymmetry

▶ EW-scale Baryogenesis

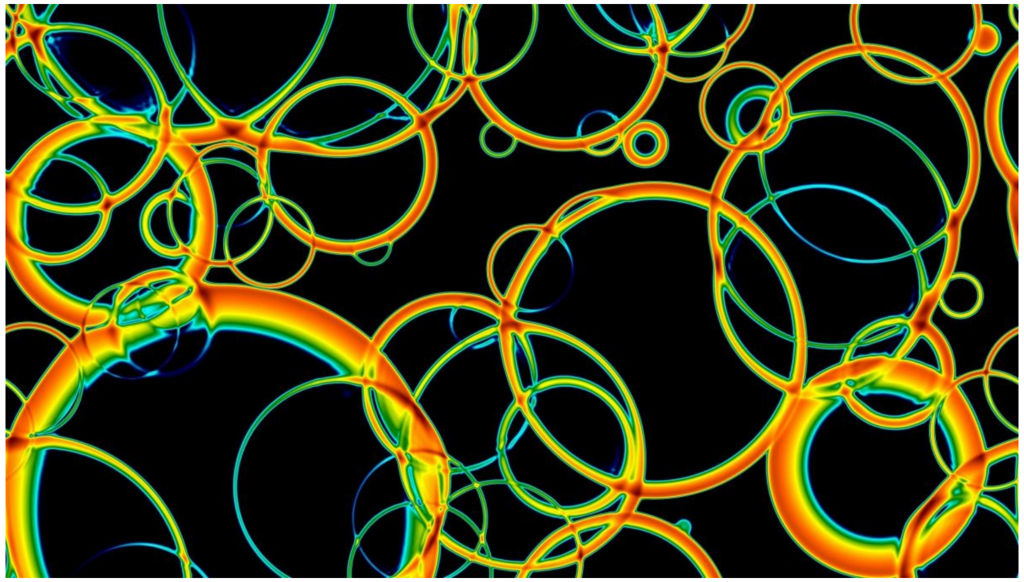
The EW Phase Transition

WHY?

- ➔ Yield Precise Understanding of EWSB in Early Universe
- ➔ (Possible) Answer to Open Mysteries at Interface of Particle Physics & Cosmology
- ➔ (Possible) Cosmological Relics from the EW Epoch

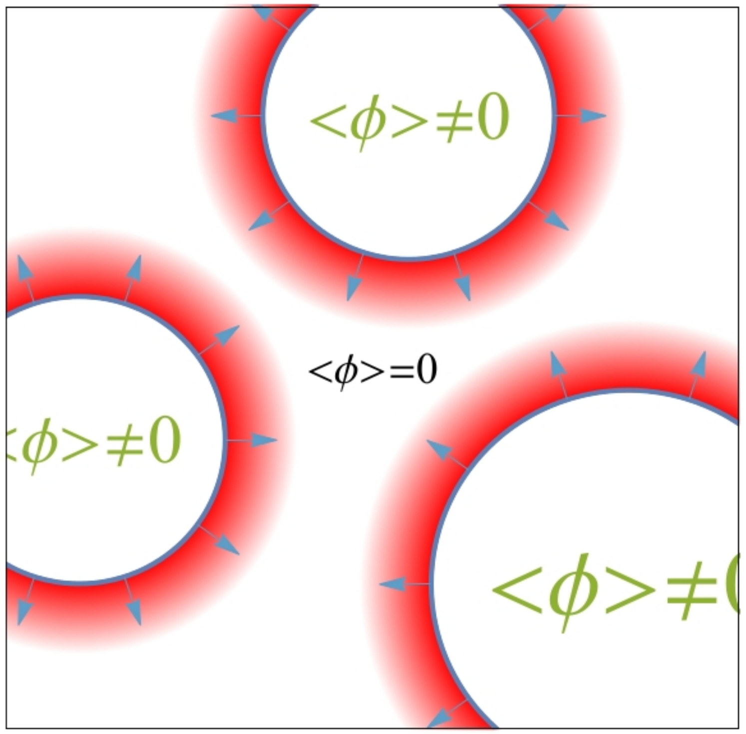
Gravitational Wave Signal

Sourced by Collisions of Higgs bubbles from a first order EW phase transition & subsequent plasma motions



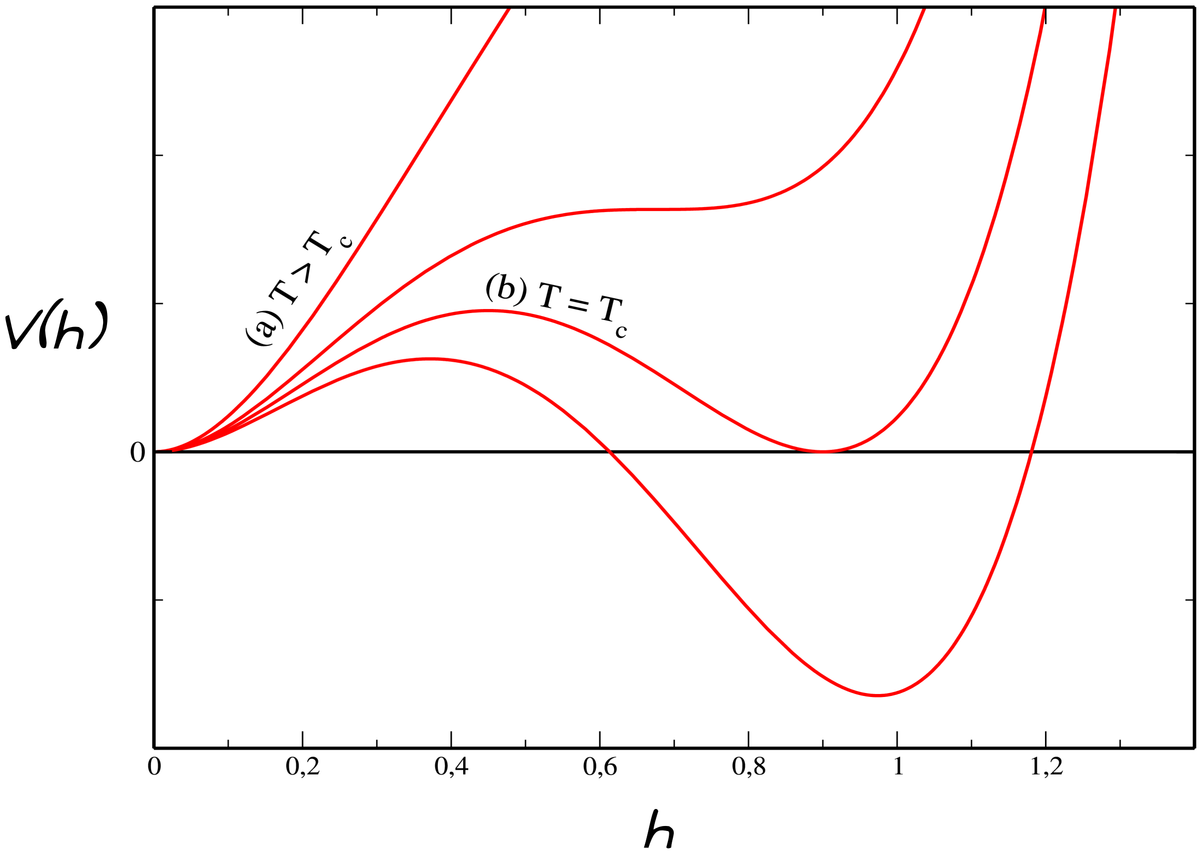
Courtesy of D. Weir (Helsinki)

*Hindmarsh, Huber, Rummukainen, Weir, PRD **92** (2015) 123009*



Assume a 1st Order EW Phase Transition…

Higgs Effective Potential (finite T)



- Phases separated by potential barrier
- Broken phase bubbles nucleate, expand, merge

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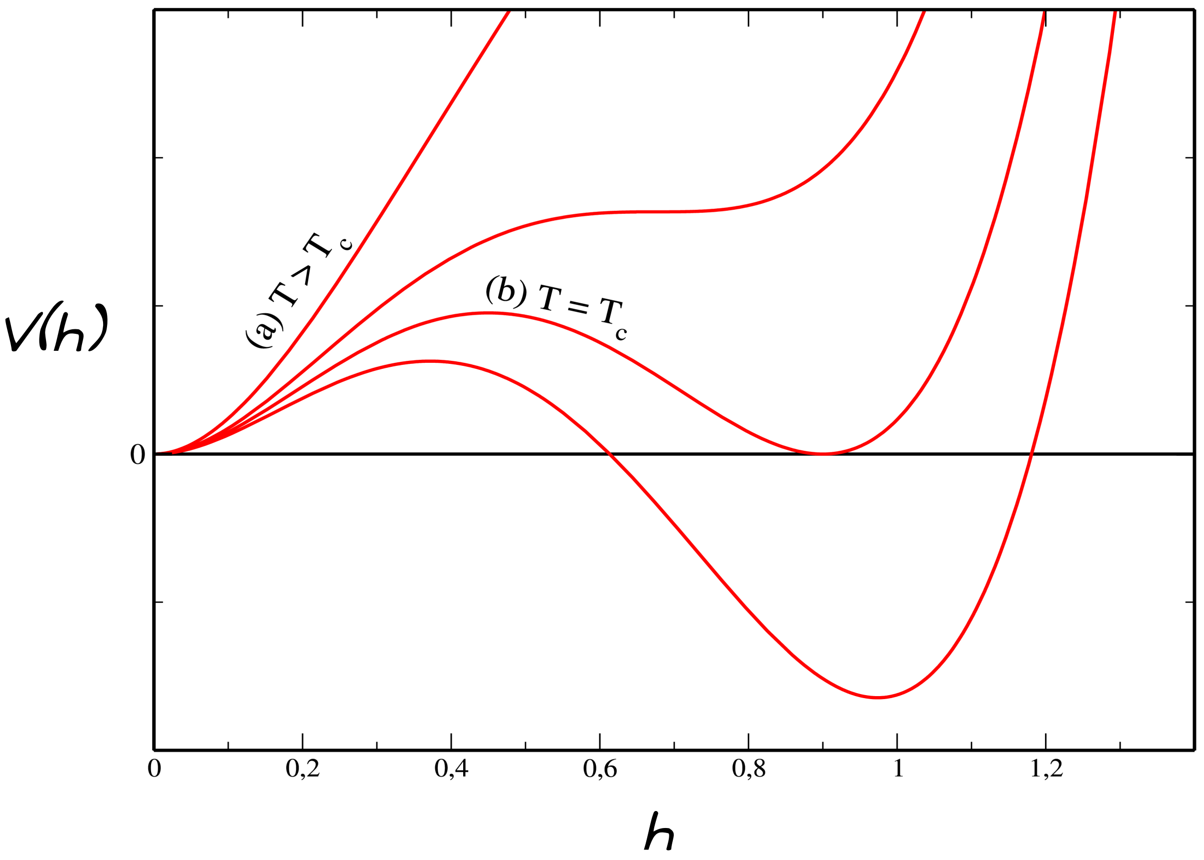

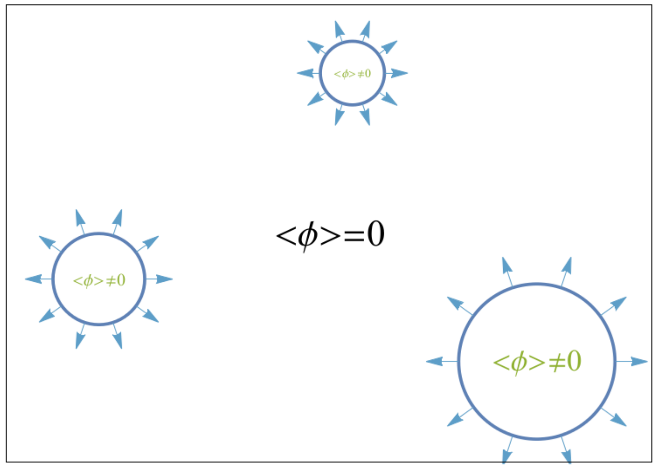

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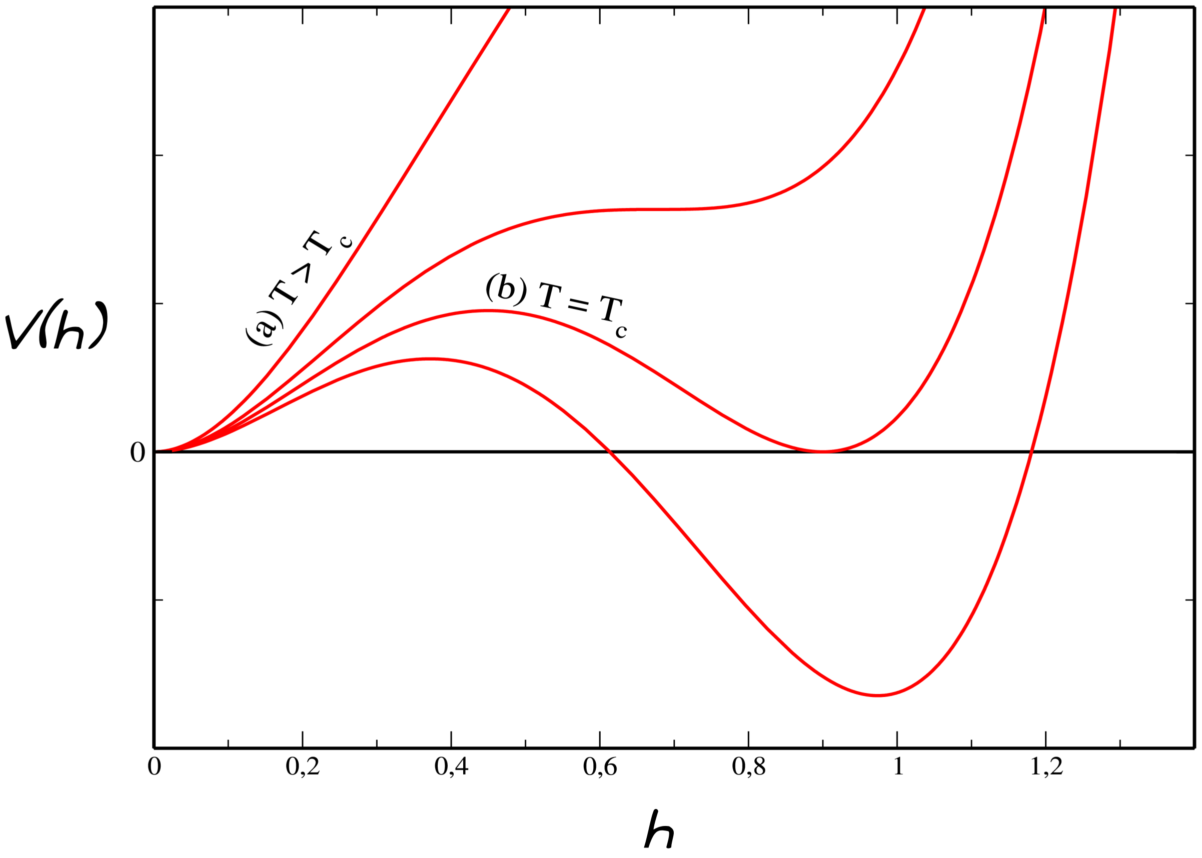
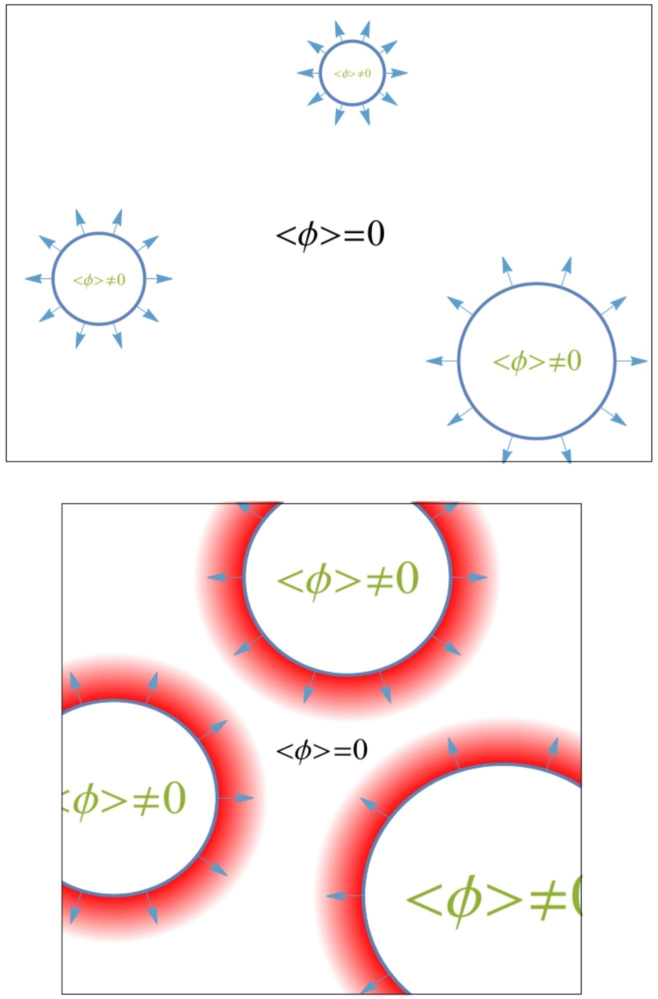
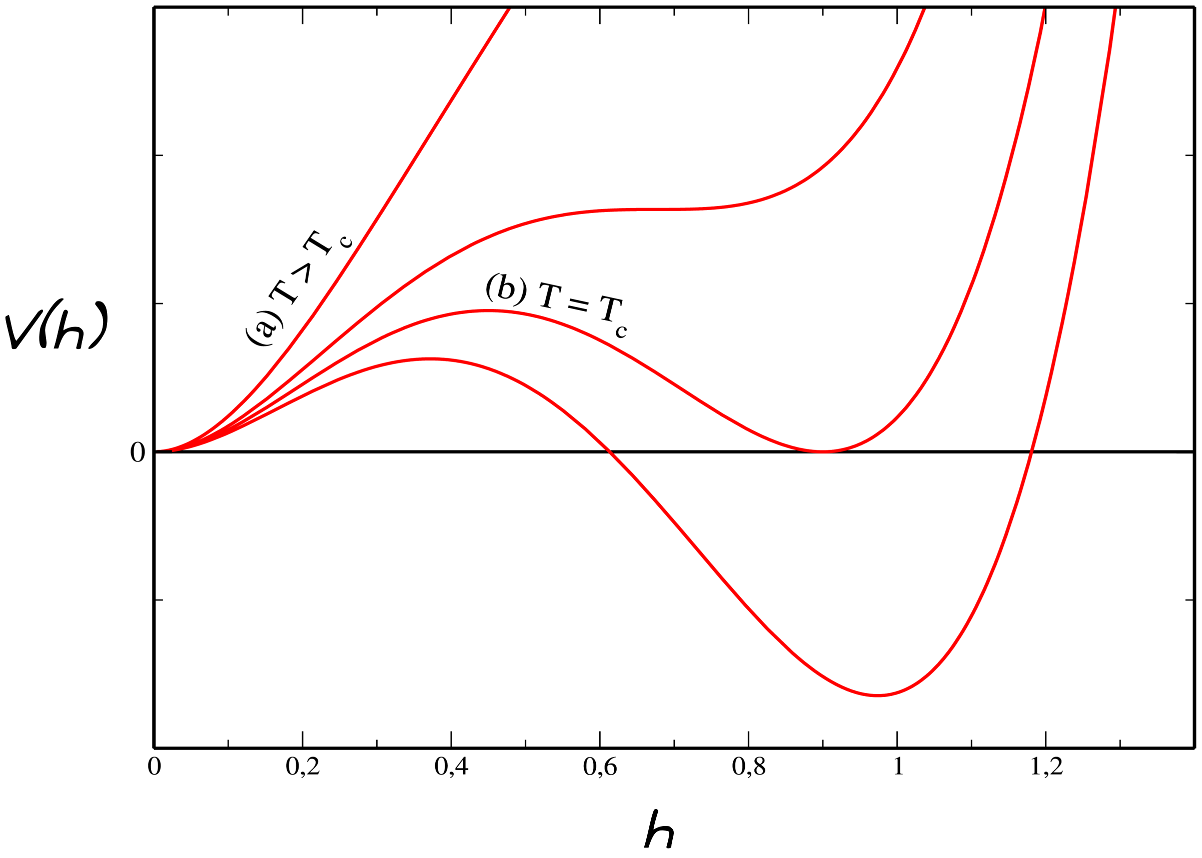




(if in plasma → create fluid waves)

- Phases separated by potential barrier
- Broken phase **bubbles** nucleate, **expand**, merge

Assume a 1st Order EW Phase Transition…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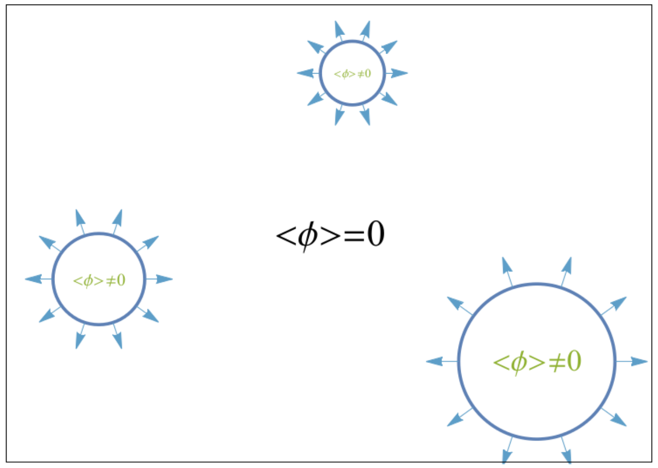

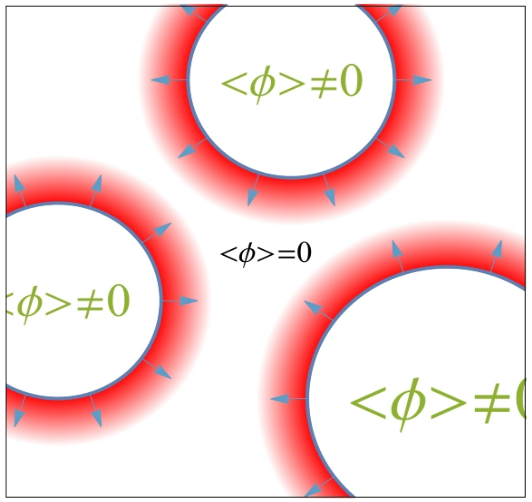

- Phases separated by potential barrier
- Broken phase **bubbles** nucleate, expand, **merge/collide** → Anisotropic Stress

⇩

Sources Gravitational Wave Production

Assume a 1st Order EW Phase Transition...

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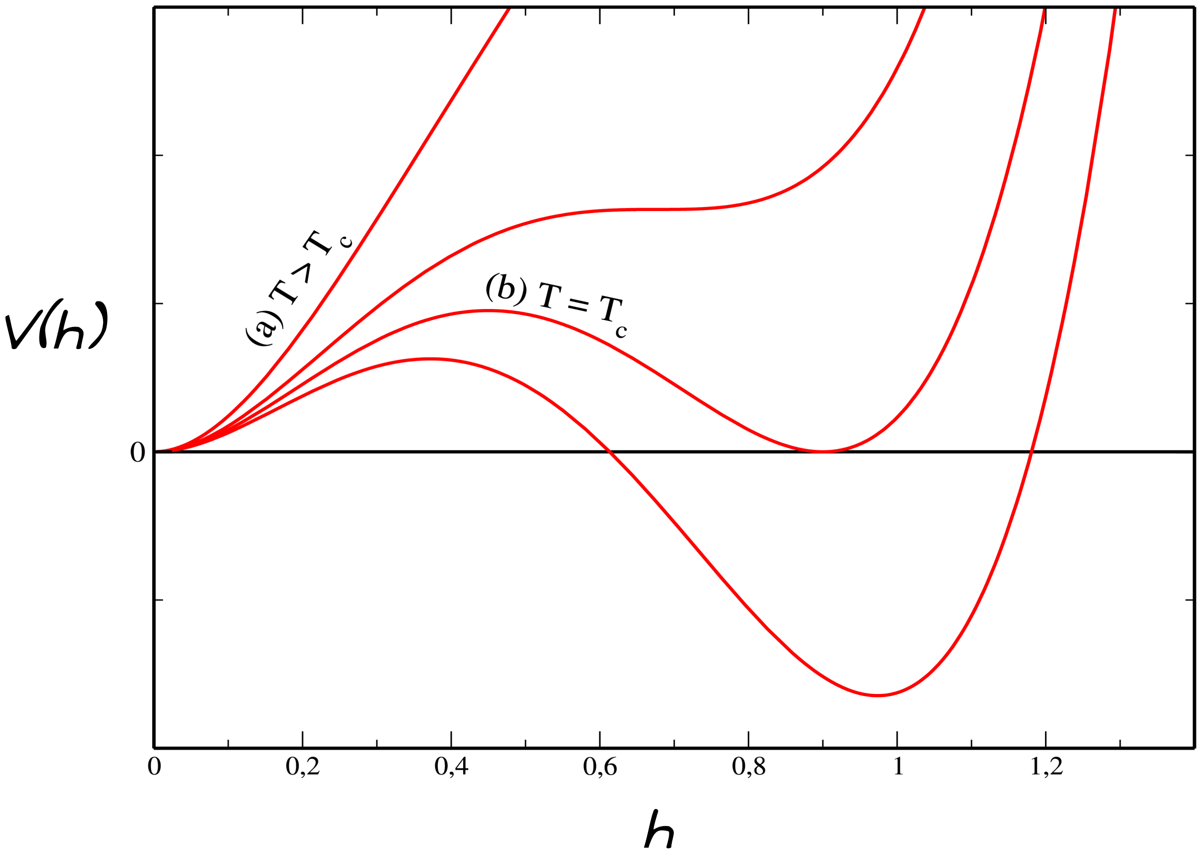
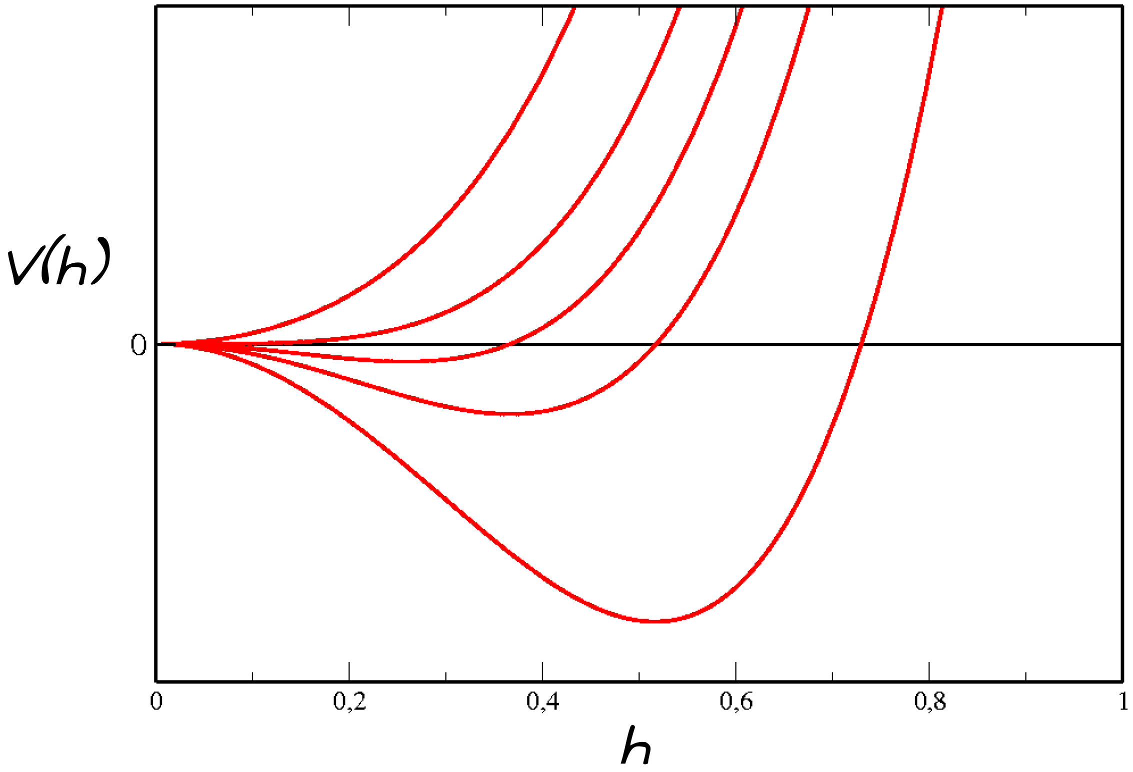


In the Standard Model...

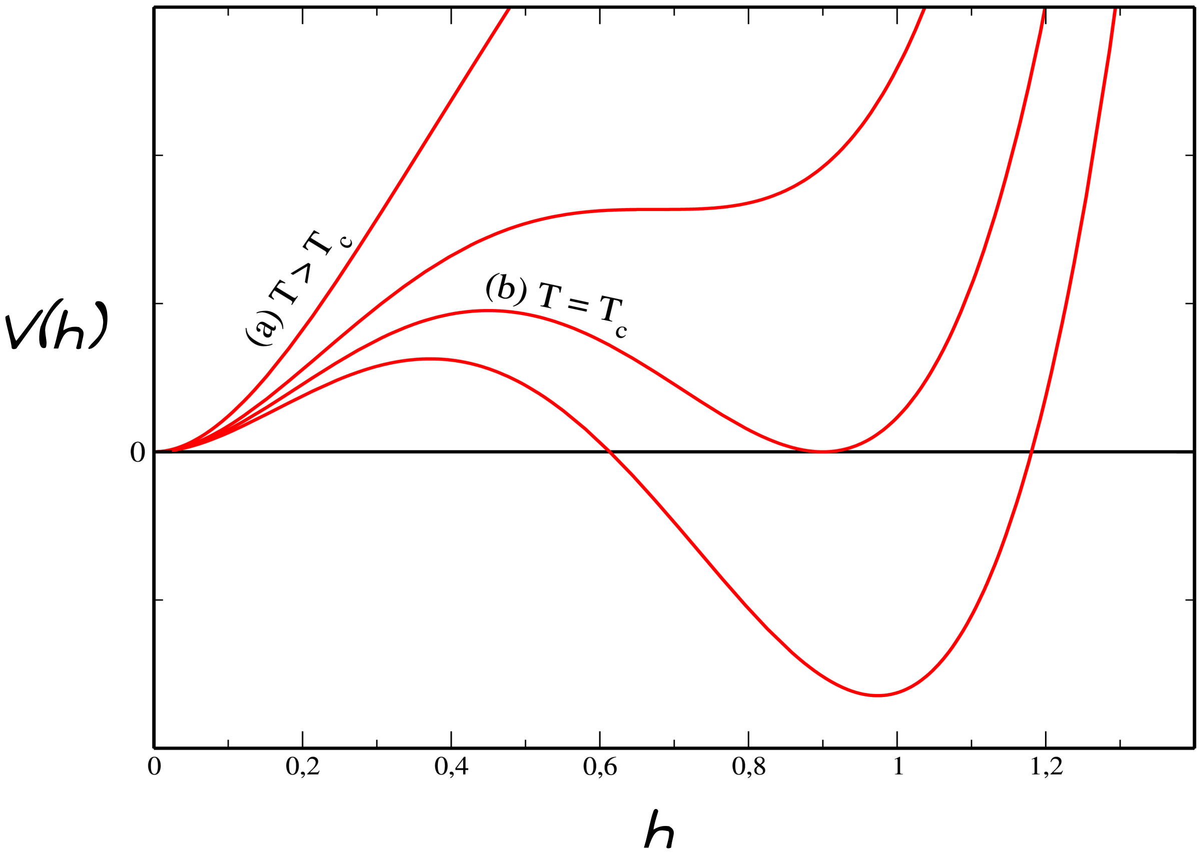


NO 1st Order EW Phase Transition

Physics Beyond the SM can induce a 1st Order EW Phase Transition

Assume a 1st Order Phase Transition…

Effective Potential (finite T)



Two “Types” of Cosmological 1st Order PTs

- “Vacuum” Transitions

Fluid/plasma effects negligible

(either plasma is very diluted or coupling between transition field and plasma small/non-existent)

Bubble walls accelerate until collision

Energy of PT stored in bubble walls

- Thermal Transitions

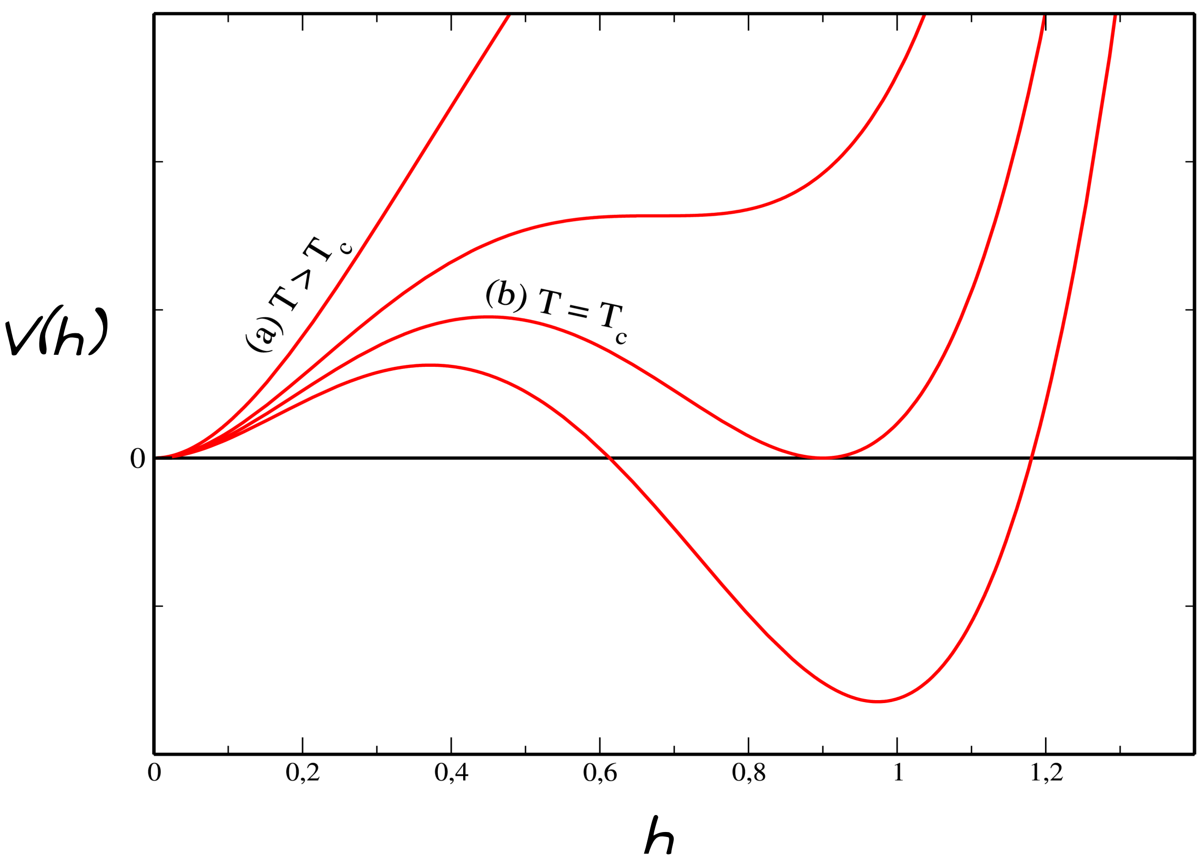
Energy of PT transferred to plasma

Plasma exerts friction on bubble wall

Terminal bubble wall velocity
(steady state)

Assume a 1st Order Phase Transition...

Effective Potential (finite T)



- Decay rate $\Gamma(T) \approx T^4 \exp\left(-\frac{S_3(T)}{T}\right)$

- O(3) symmetric action

$$S_3(T) = 4\pi \int dr r^2 \left[\frac{1}{2}\left(\frac{d\phi}{dr}\right)^2 + V(\phi, T)\right]$$

- Bubble profile (bounce)

$$\frac{d^2\phi}{dr^2} + \frac{2}{r}\frac{d\phi}{dr} - \frac{\partial V(\phi, T)}{\partial \phi} = 0$$

$$\phi(r \to \infty) = 0 \quad \text{and} \quad \dot{\phi}(r = 0) = 0$$

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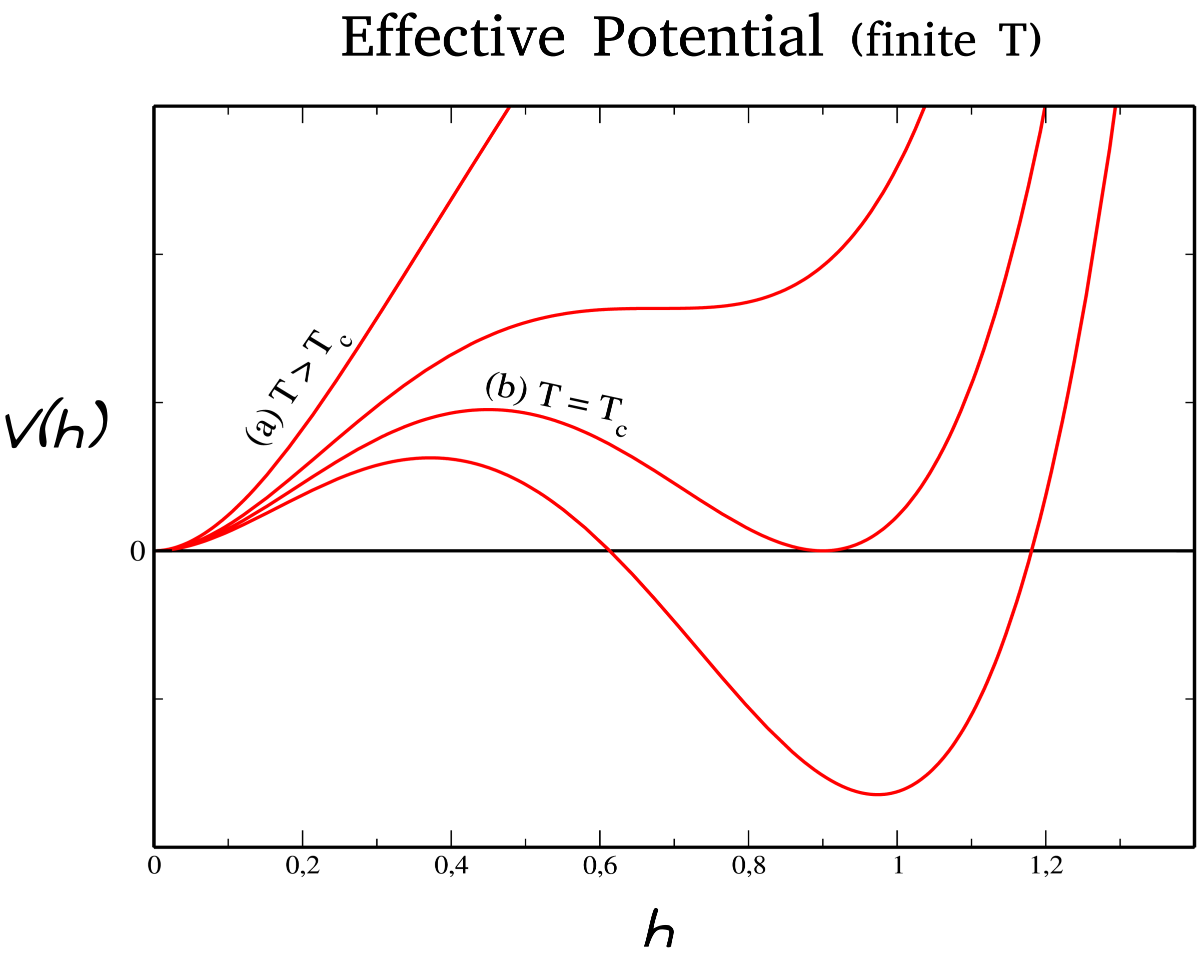
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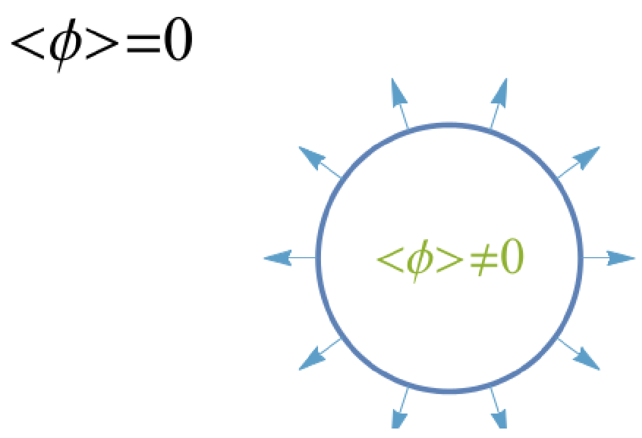

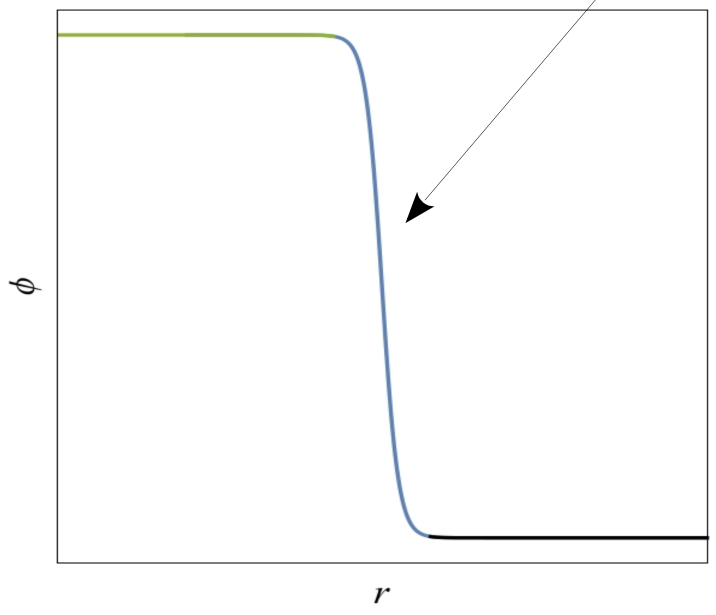

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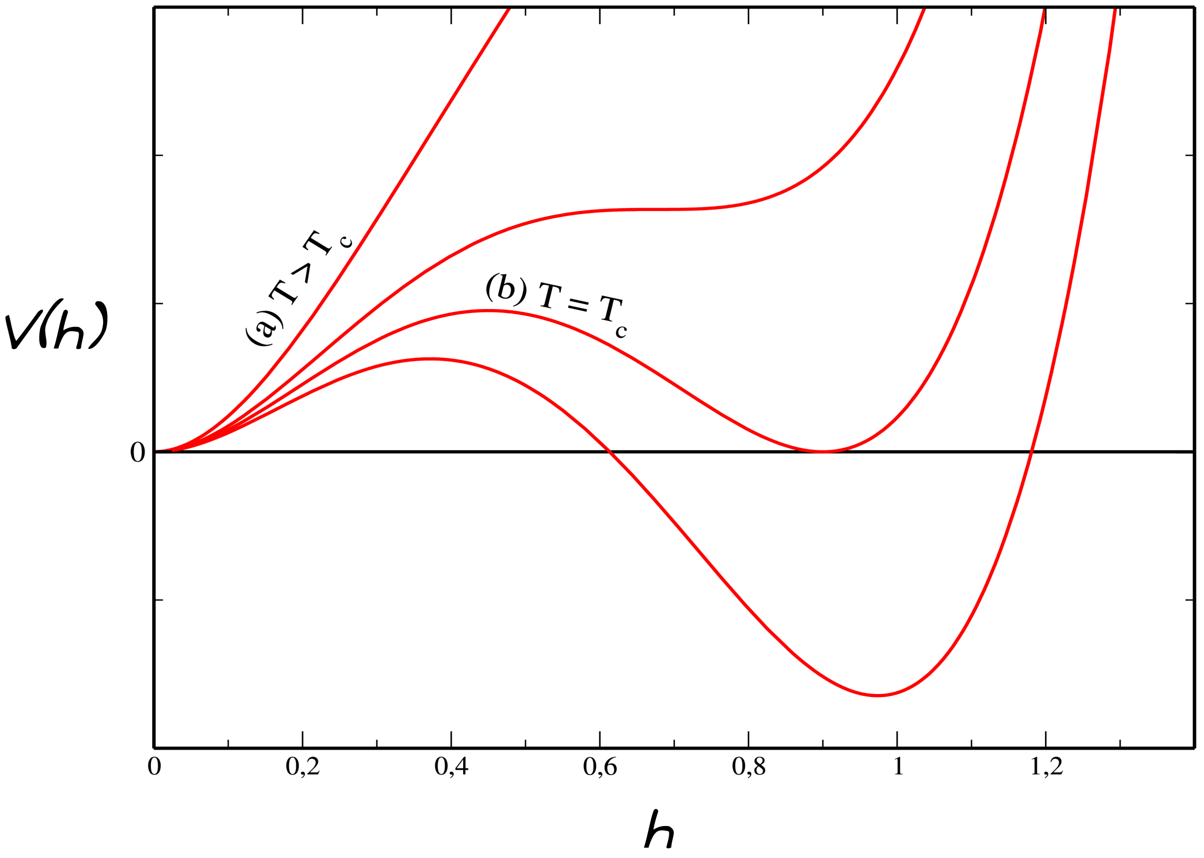

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Nucleation temperature:

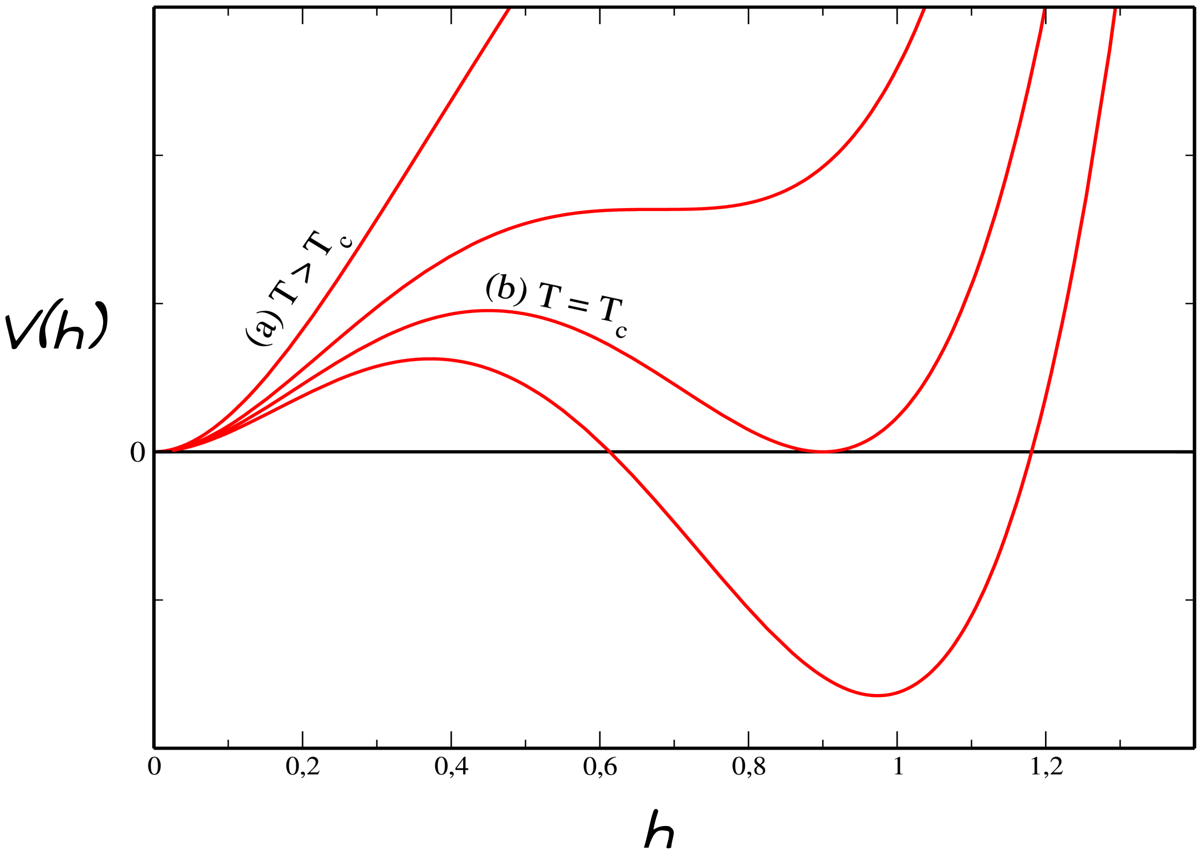
One Higgs bubble per Horizon volume (on average)

$$N(T_n) = \int_{t_c}^{t_n} dt \frac{\Gamma(t)}{H(t)^3} = \int_{T_n}^{T_c} \frac{dT}{T} \frac{\Gamma(T)}{H(T)^4} = 1$$

*Linde, Phys. Lett. B**100** (1981) 37; Nucl. Phys. B**216** (1983) 421*

Assume a 1st Order Phase Transition...

Effective Potential (finite T)



Two KEY Phase Transition Quantities:

- (Available) Transition Energy (normalized)

$$\alpha = \frac{\epsilon}{a_+ T^4}$$

Assume a 1st Order Phase Transition...

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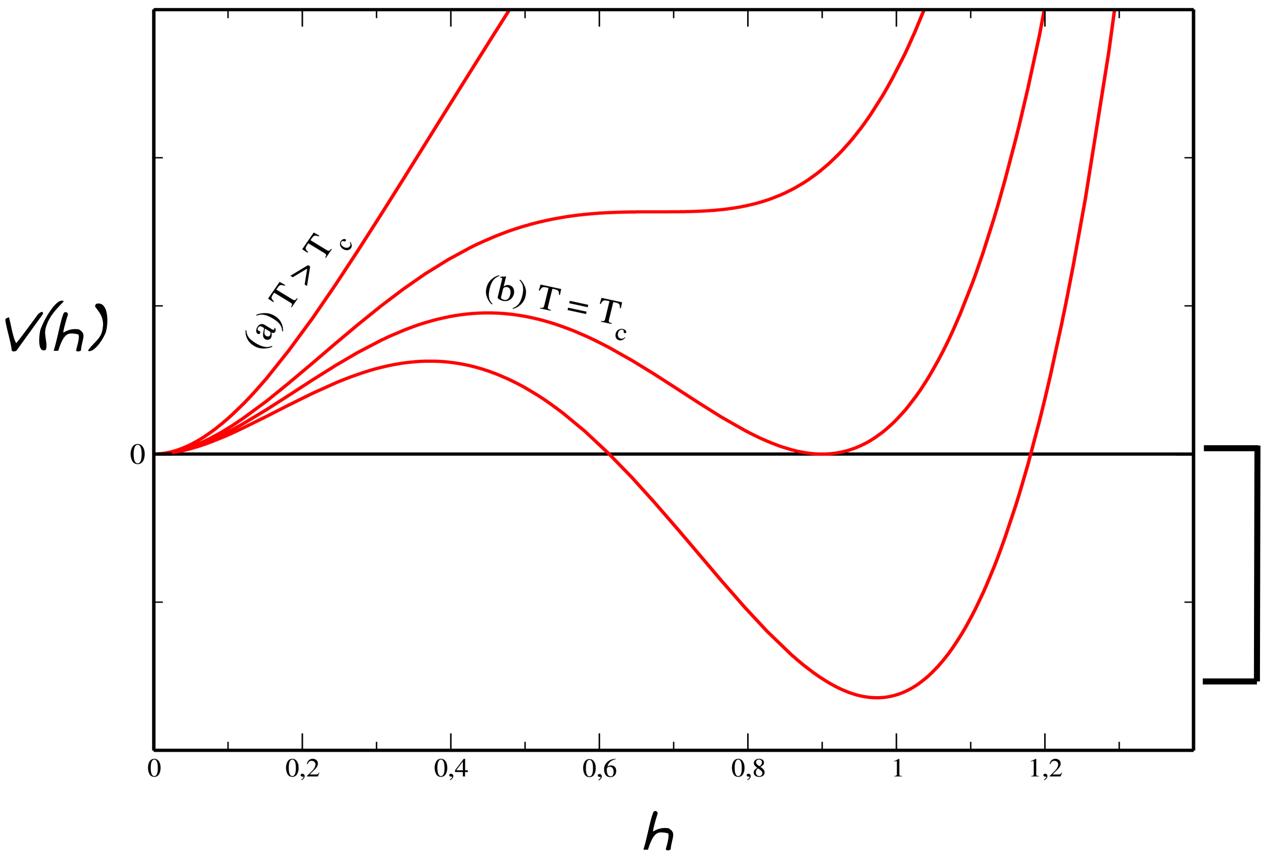

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Radiation Energy Density

Assume a 1st Order Phase Transition…

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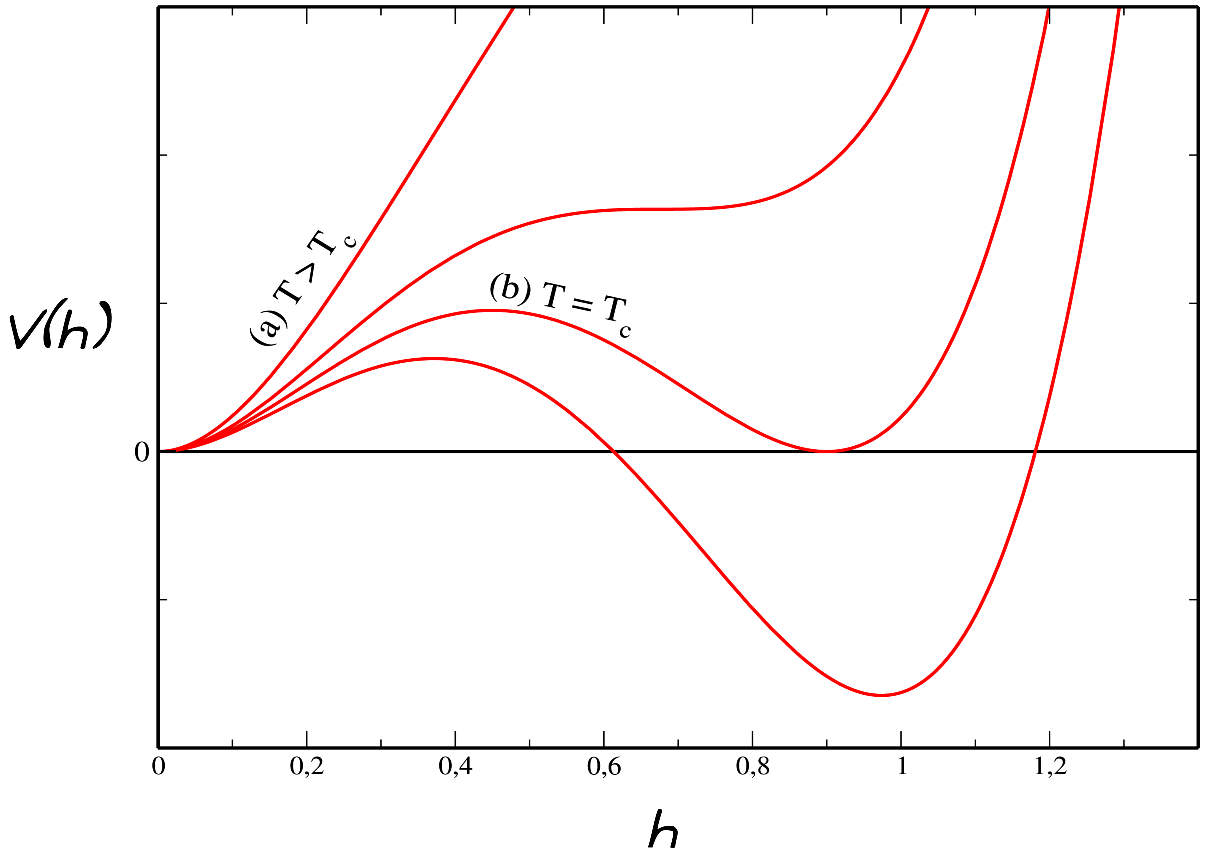

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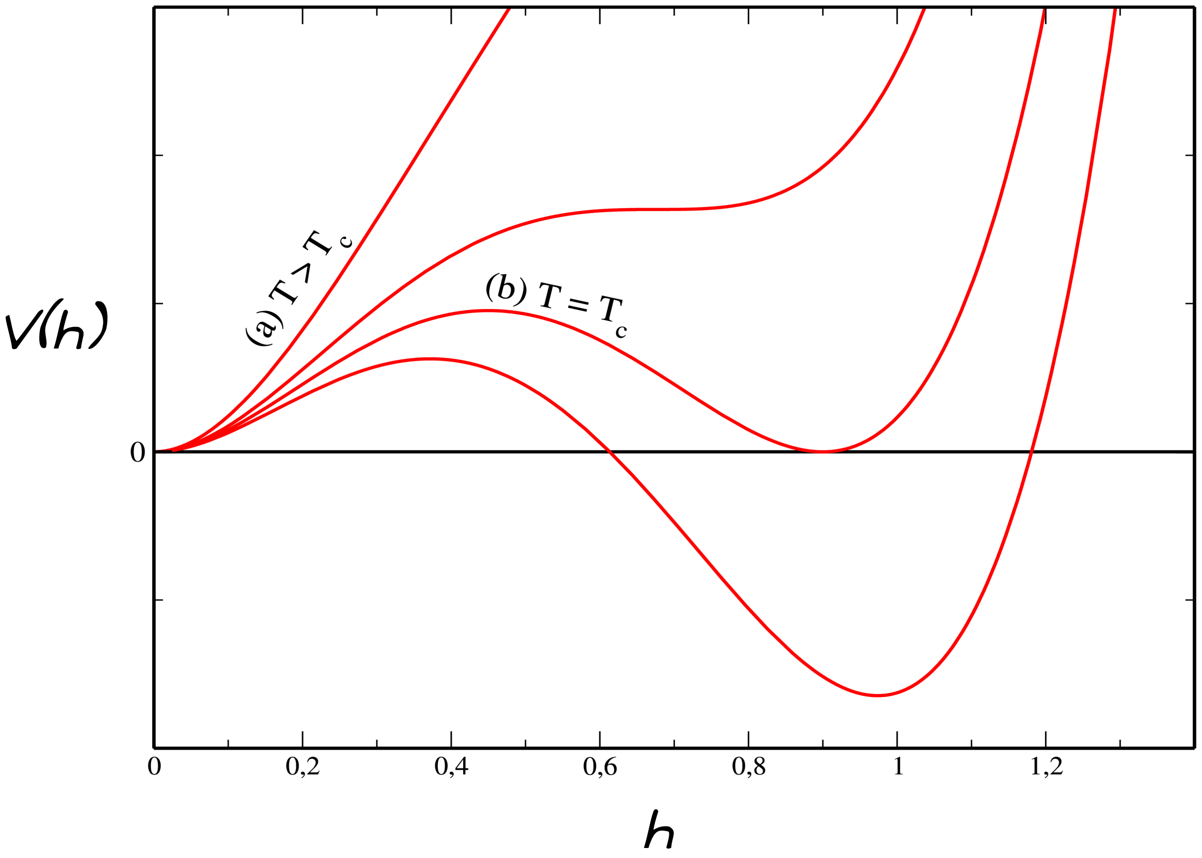
Assume a 1^{st} Order Phase Transition…

Effective Potential (finite T)



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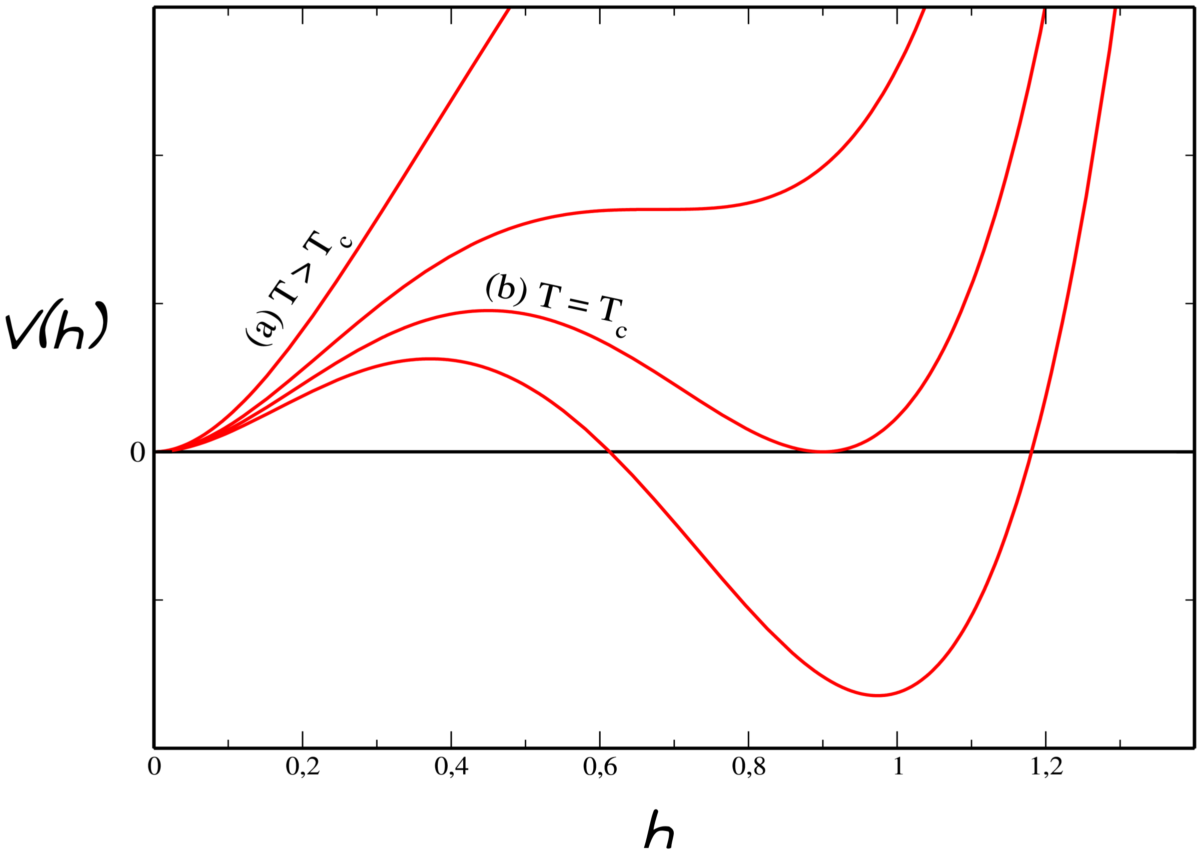
$$\frac{\beta}{H} \equiv -\frac{dS_3}{dt}\Big|_{t=t_n} \approx T\frac{d(S_3/T)}{dT}\Big|_{T=T_n}$$

(Related to the change of the Decay Rate)

Assume a 1st Order Phase Transition...

Effective Potential (finite T)



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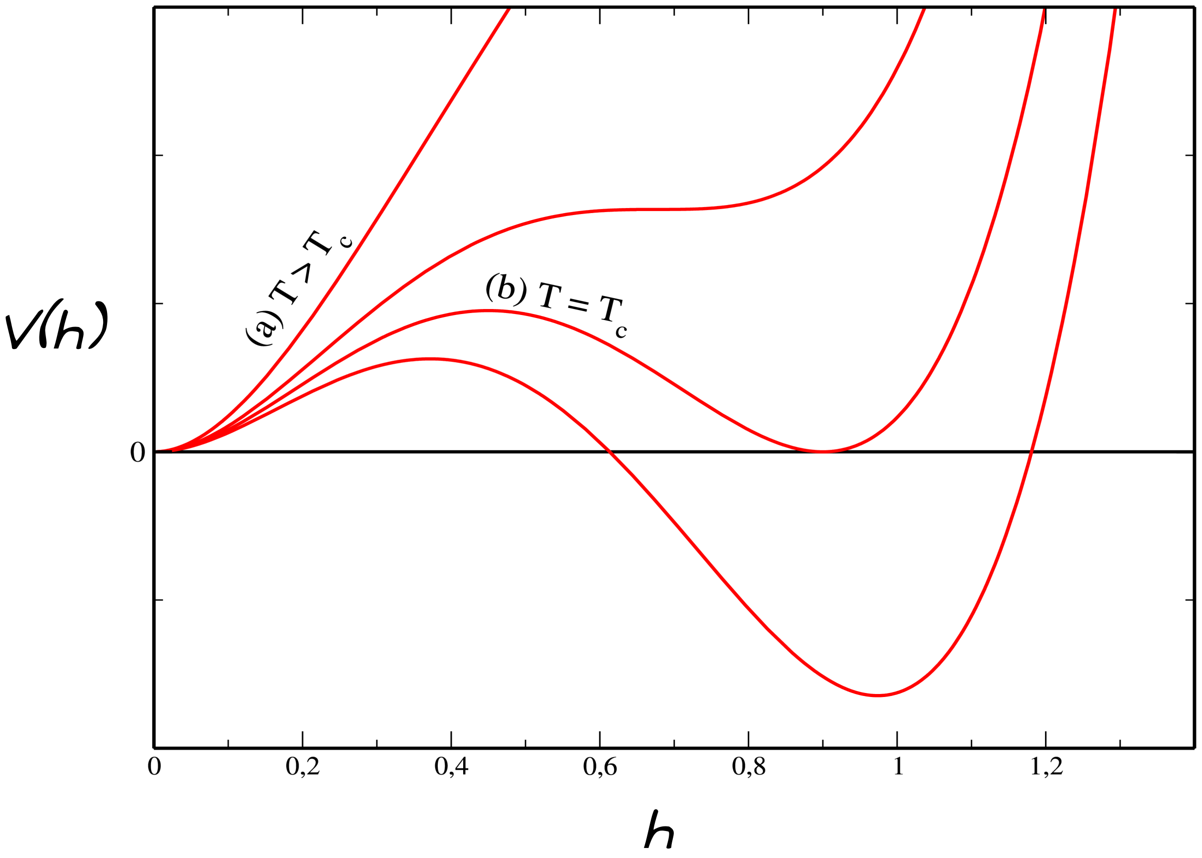
Average number of bubbles per horizon at the time of bubble coalescence/percolation

(Transition Completes, T_* **)**

H_*

Assume a 1^{st} Order Phase Transition...

Effective Potential (finite T)



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(Related to the change of the Decay Rate)

Average number of bubbles per horizon at the time of bubble coalescence/percolation

(Transition Completes, T_*)

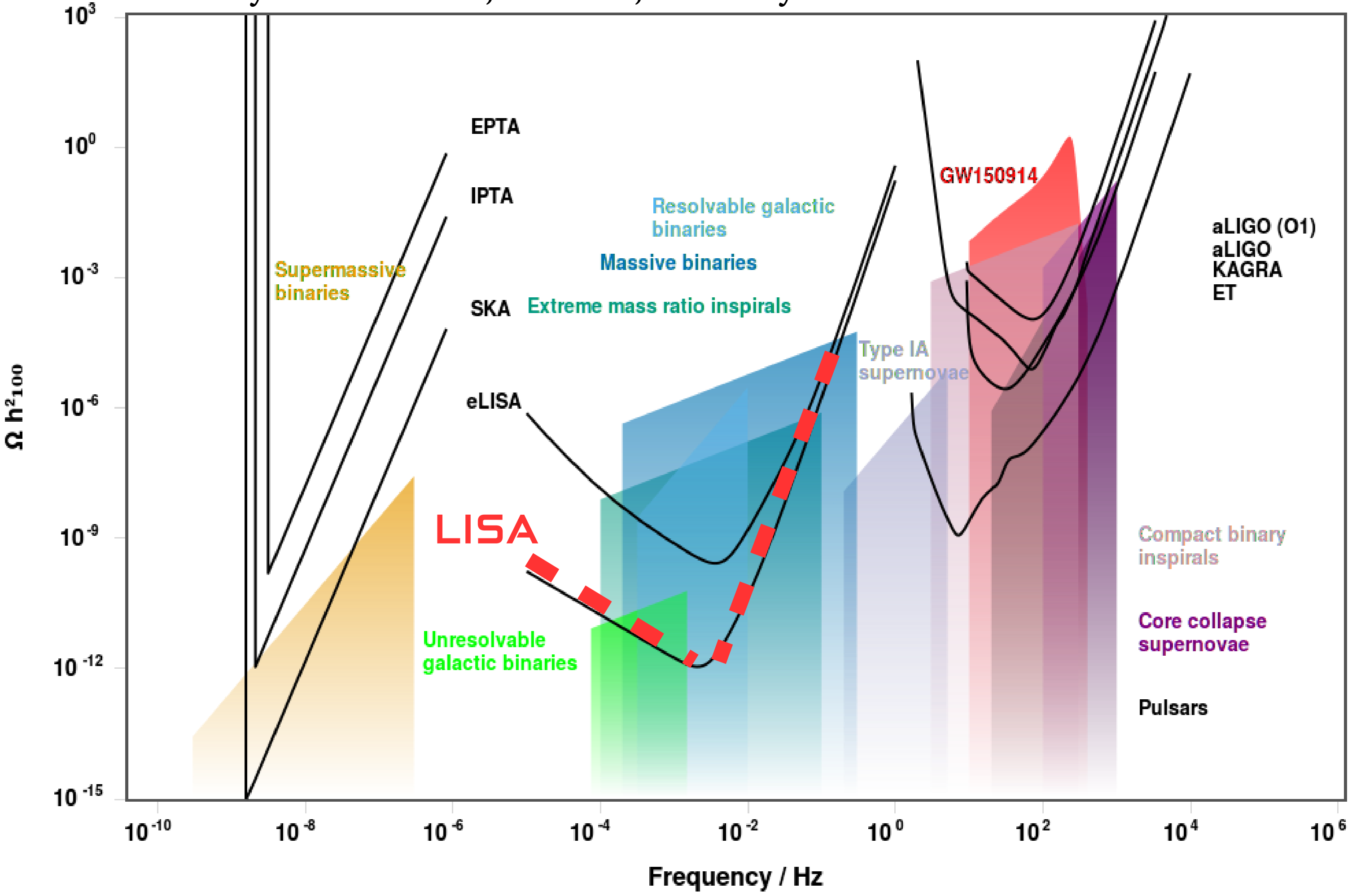
H_*

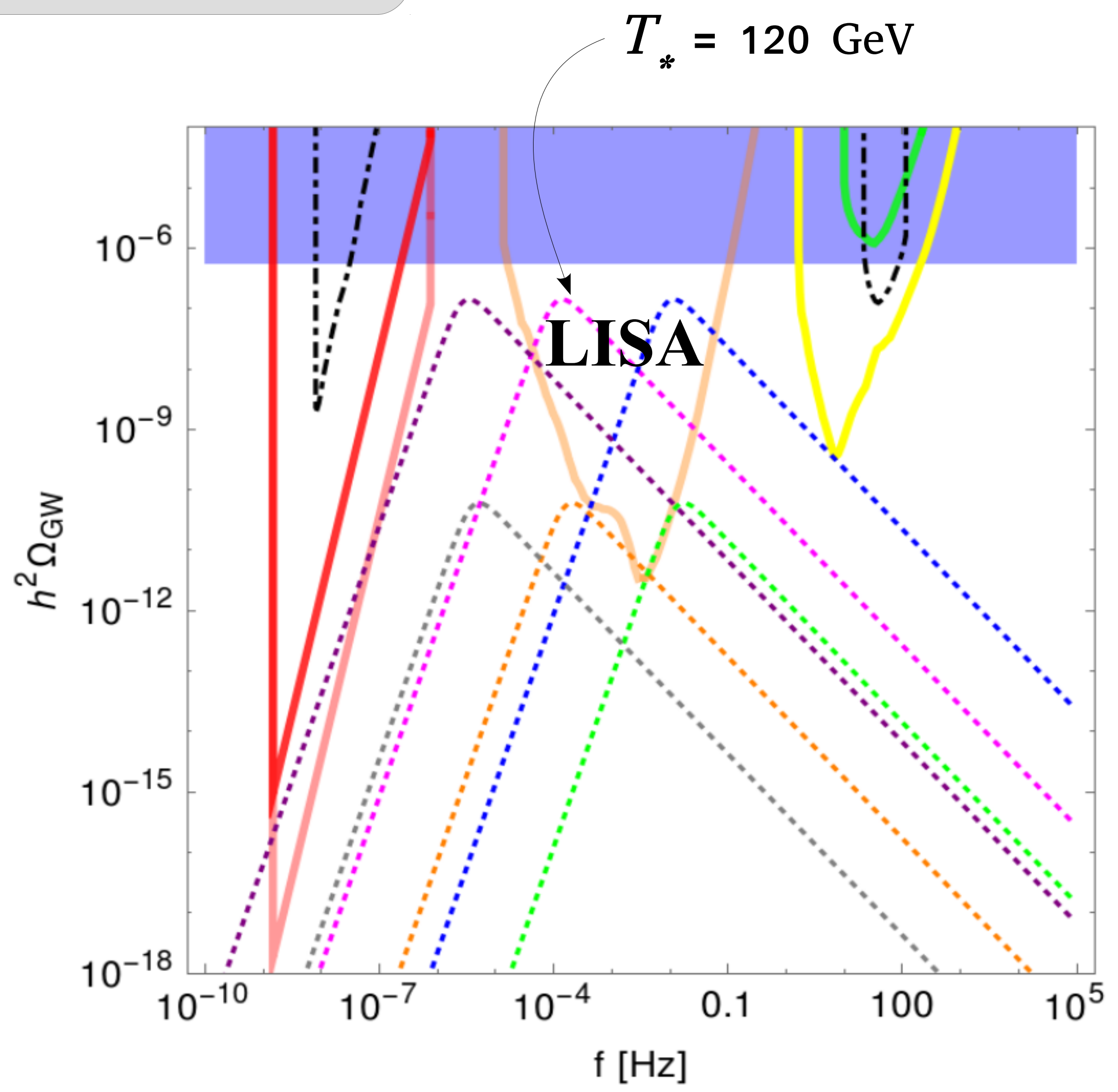
- GW frequency ~ **size of bubbles @ collision**
- For $T_* \sim 100$ GeV and $\frac{\beta}{H_*} \sim 100$, GW frequency (redshifted to today!) ~ **mHz**

1^{st} Order (EW) Phase Transition

mHz GW Signal in the sensitivity band of future space-based GW detector **LISA**

Courtesy of C. Moore, R. Cole, C. Berry (GWPlotter)





$T_* = 120$ GeV

Figueroa et al., PoS GRASS2018 (2018) 036

1^{st} Order (EW) Phase Transition

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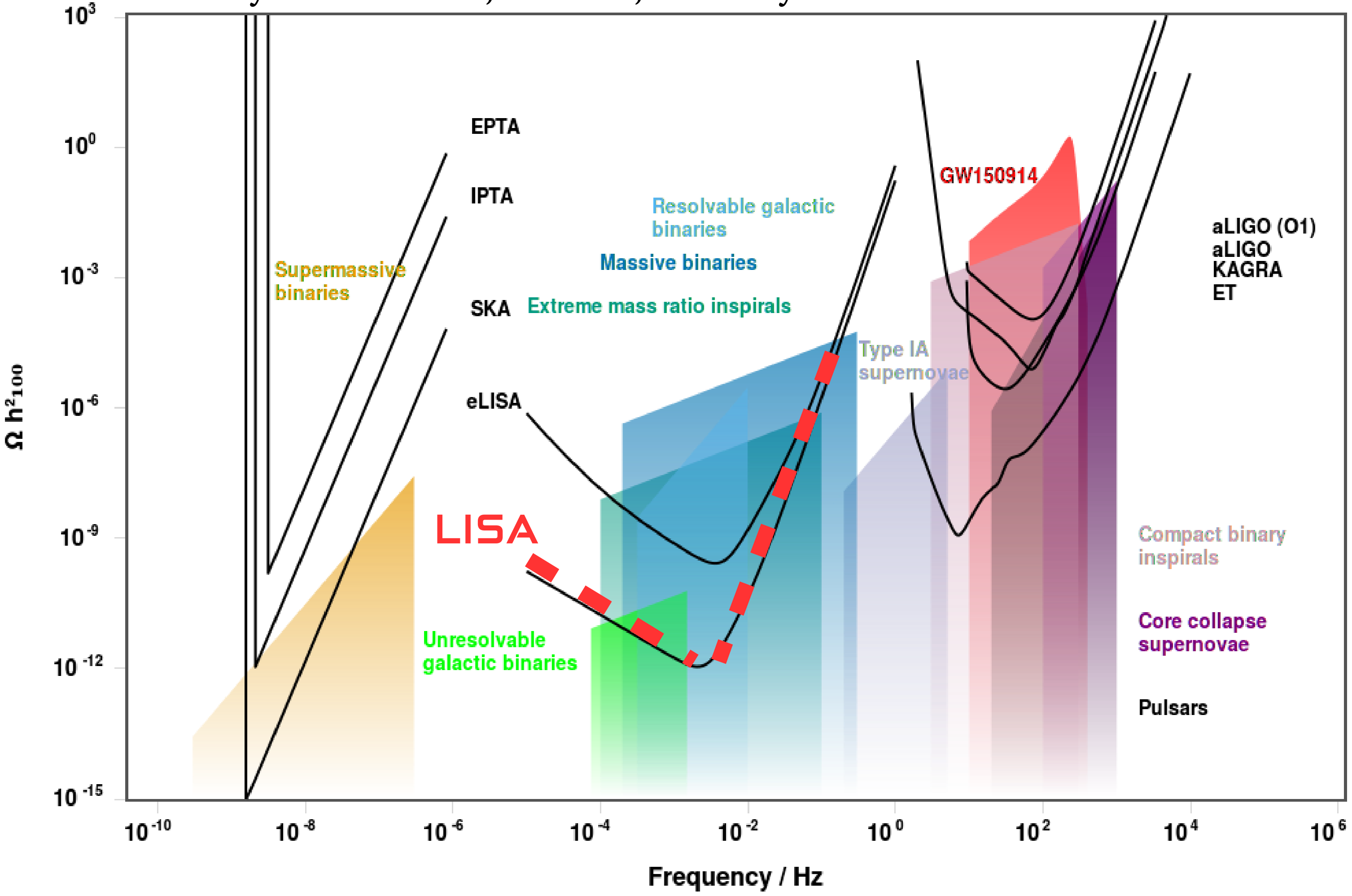
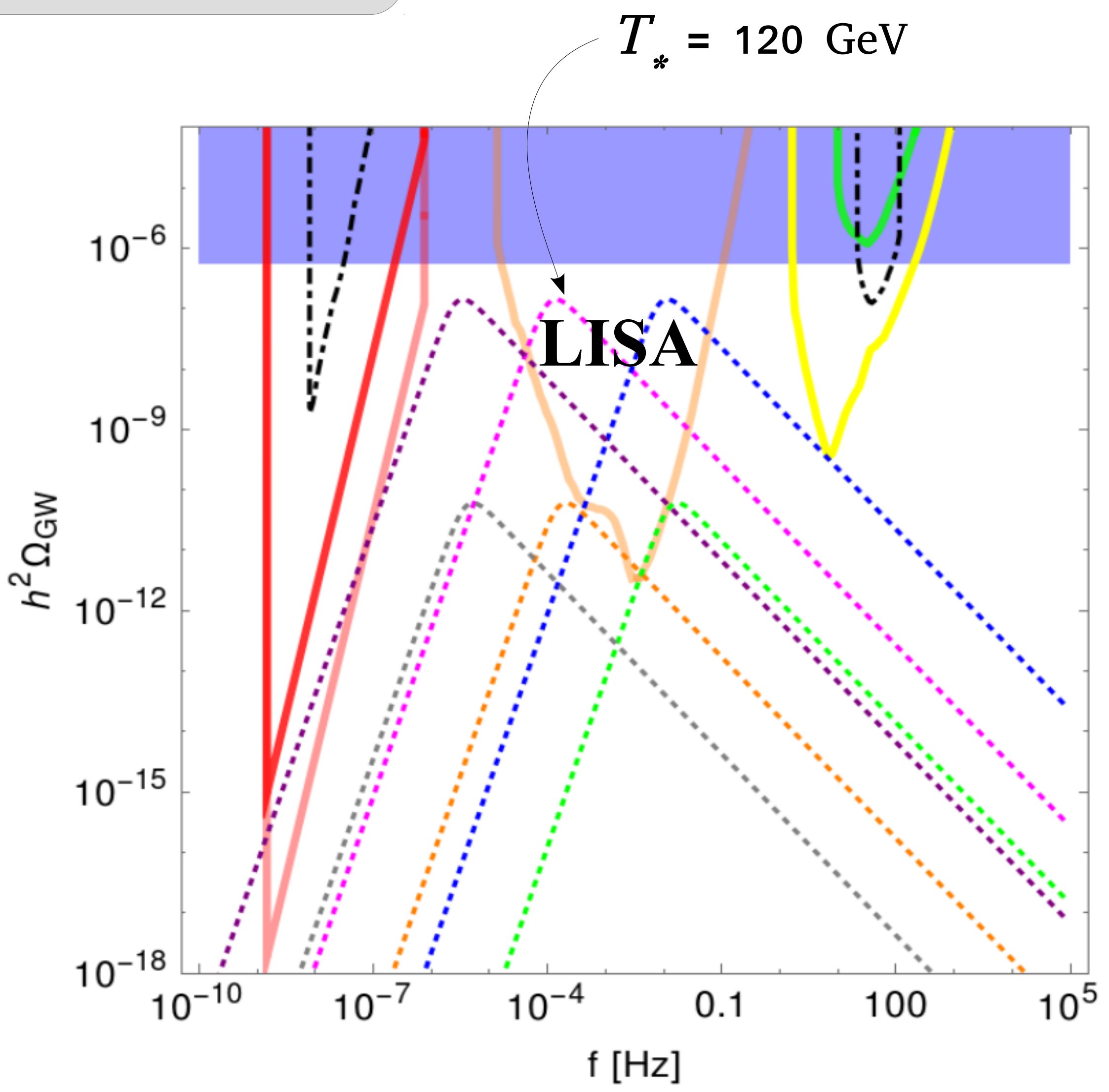




Figueroa et al., PoS GRASS2018 (2018) 036

LISA can probe the **EW epoch** of the early Universe

(

https://www.elisascience.org/

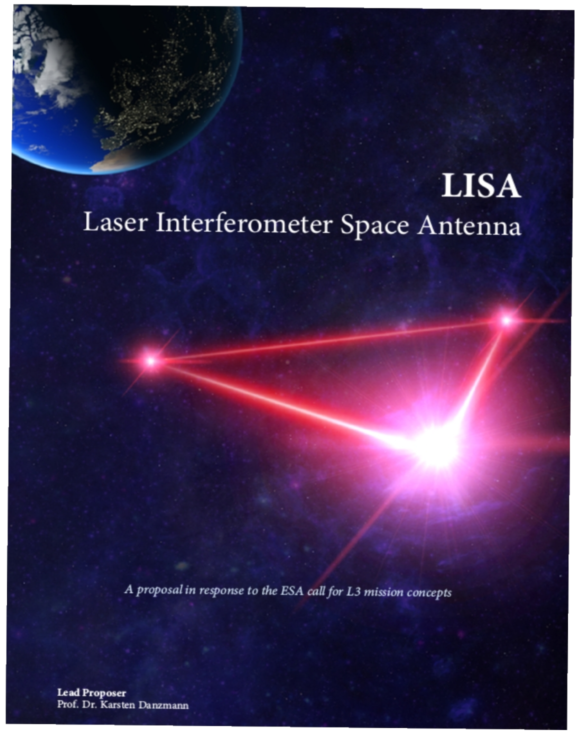
The LISA Mission

(Laser Interferometer Space Antenna)

A brief status report

Thanks to G. Nardini

2017: LISA proposal to ESA



LISA

Laser Interferometer Space Antenna

A proposal in response to the ESA call for L3 mission concepts

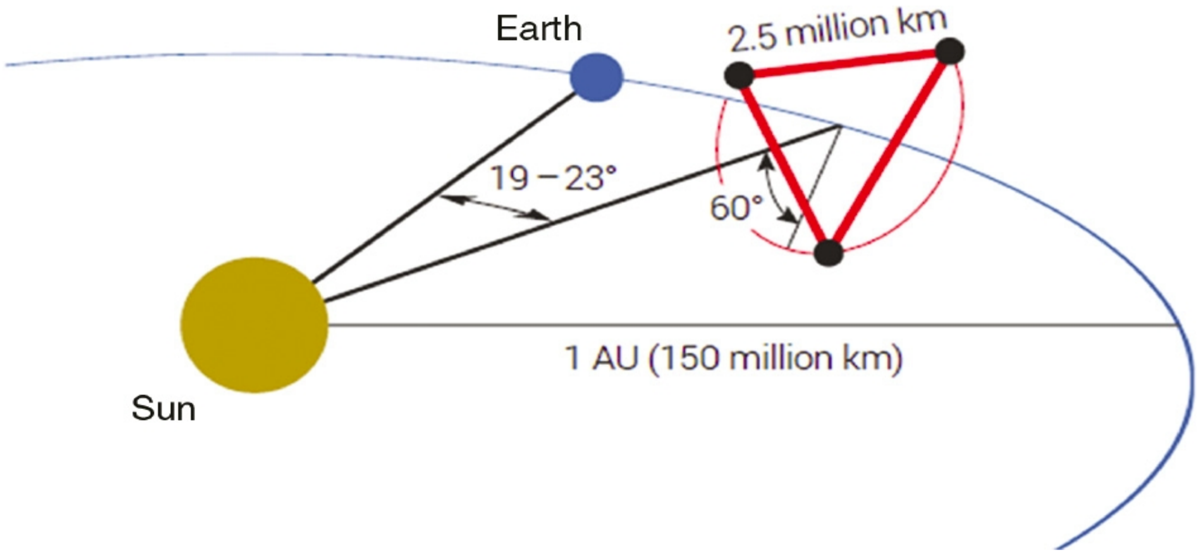
Lead Proposer
Prof. Dr. Karsten Danzmann

LISA Collaboration, 1702.00786

Launch date **2030-2034**

LISA Mission selected by ESA (Summer **2017**)

+ (On Jan 22 2018, LISA passed ESA's Mission Definition Review)

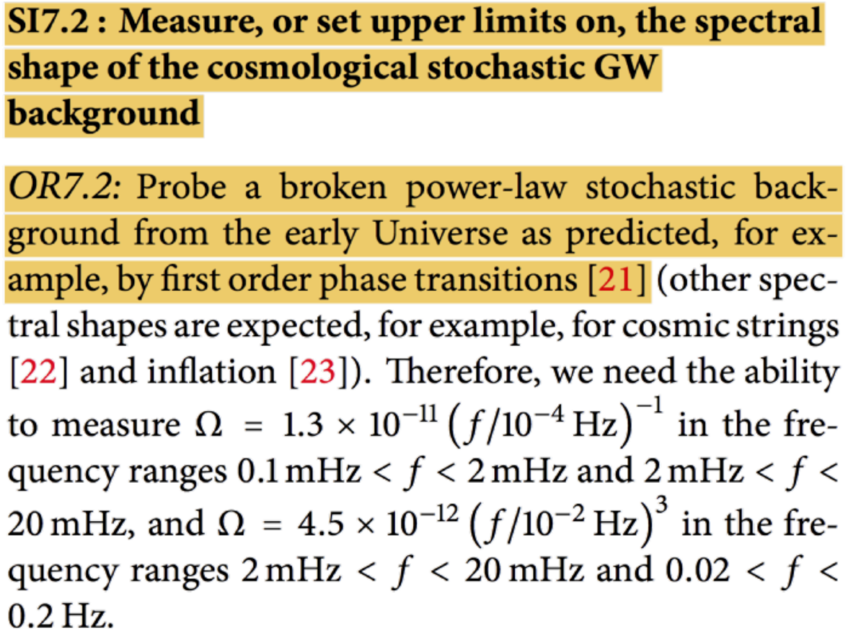


4 years of lifetime (w. consumables up to 10 years)

2.5 MKm (arm length)

From the proposal:

Audley et al, arXiv:1702.00786

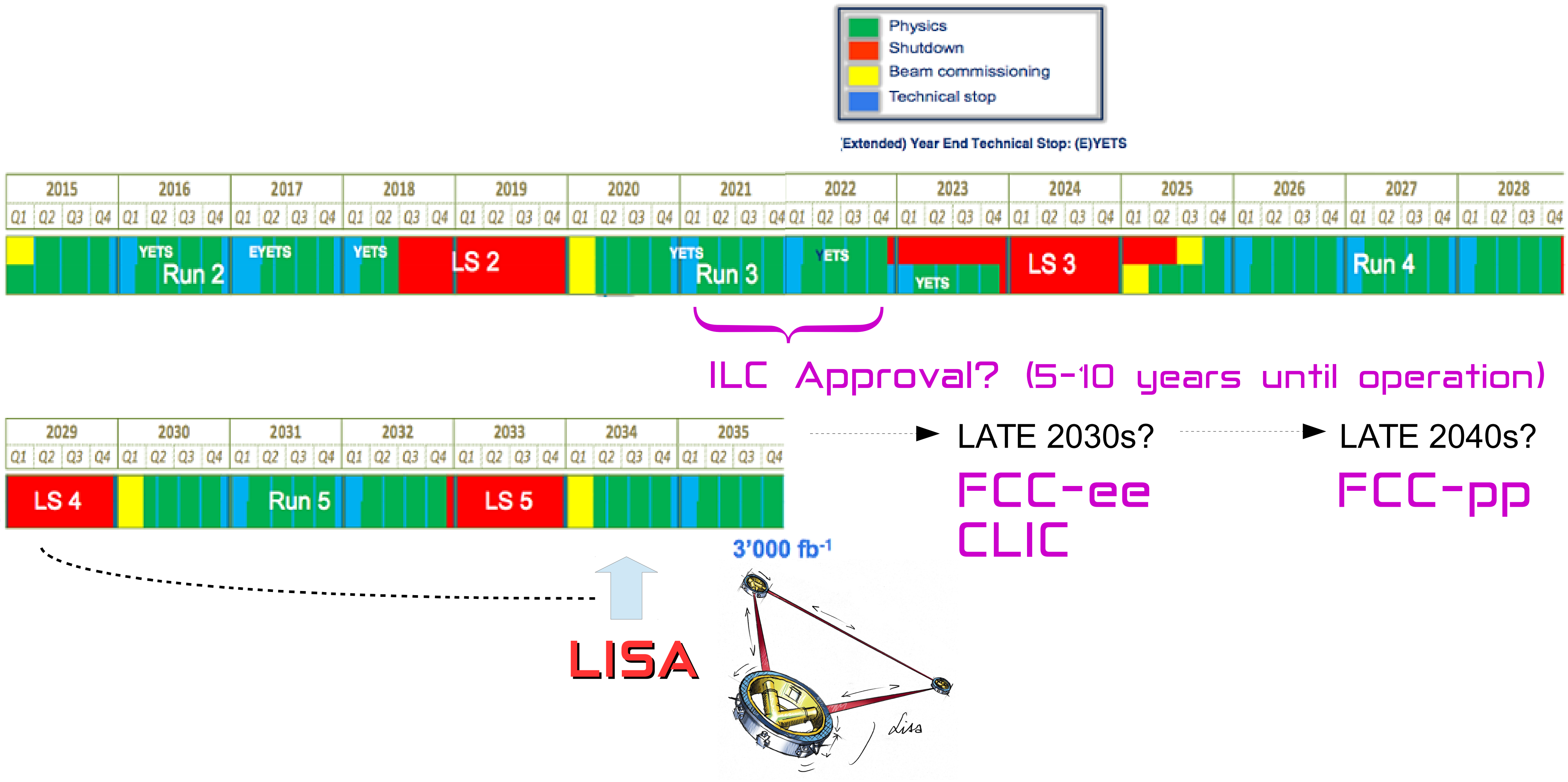


SI7.2 : Measure, or set upper limits on, the spectral shape of the cosmological stochastic GW background

OR7.2: Probe a broken power-law stochastic background from the early Universe as predicted, for example, by first order phase transitions [21] (other spectral shapes are expected, for example, for cosmic strings [22] and inflation [23]). Therefore, we need the ability to measure $\Omega = 1.3 \times 10^{-11} \left(f/10^{-4}\,\mathrm{Hz}\right)^{-1}$ in the frequency ranges $0.1\,\mathrm{mHz} < f < 2\,\mathrm{mHz}$ and $2\,\mathrm{mHz} < f < 20\,\mathrm{mHz}$, and $\Omega = 4.5 \times 10^{-12} \left(f/10^{-2}\,\mathrm{Hz}\right)^{3}$ in the frequency ranges $2\,\mathrm{mHz} < f < 20\,\mathrm{mHz}$ and $0.02 < f < 0.2\,\mathrm{Hz}$.

GW – Collider complementarity

Timeline: LISA GW Observatory in the Context of High-Energy Colliders



ILC Approval? (5-10 years until operation)

LATE 2030s?

FCC-ee

CLIC

LATE 2040s?

FCC-pp

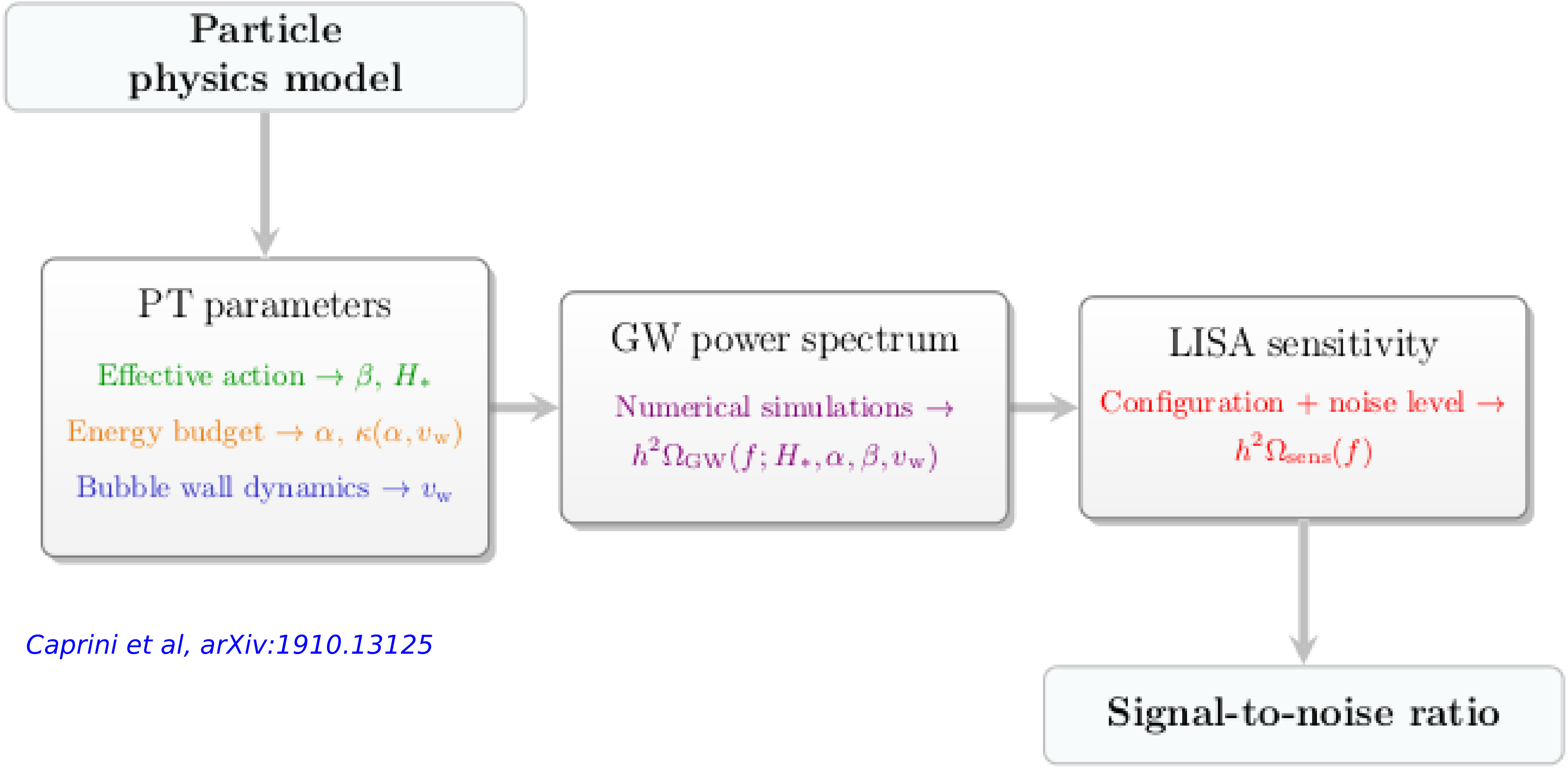
$3'000\ \mathrm{fb}^{-1}$

LISA

After LHC, LISA is next step in exploration of EW scale physics

)

GW from the EW Phase Transition with **LISA**



Particle physics model

PT parameters

Effective action → β, H_*

Energy budget → α, $\kappa(\alpha, v_w)$

Bubble wall dynamics → v_w

GW power spectrum

Numerical simulations → $h^2\Omega_{GW}(f; H_*, \alpha, \beta, v_w)$

LISA sensitivity

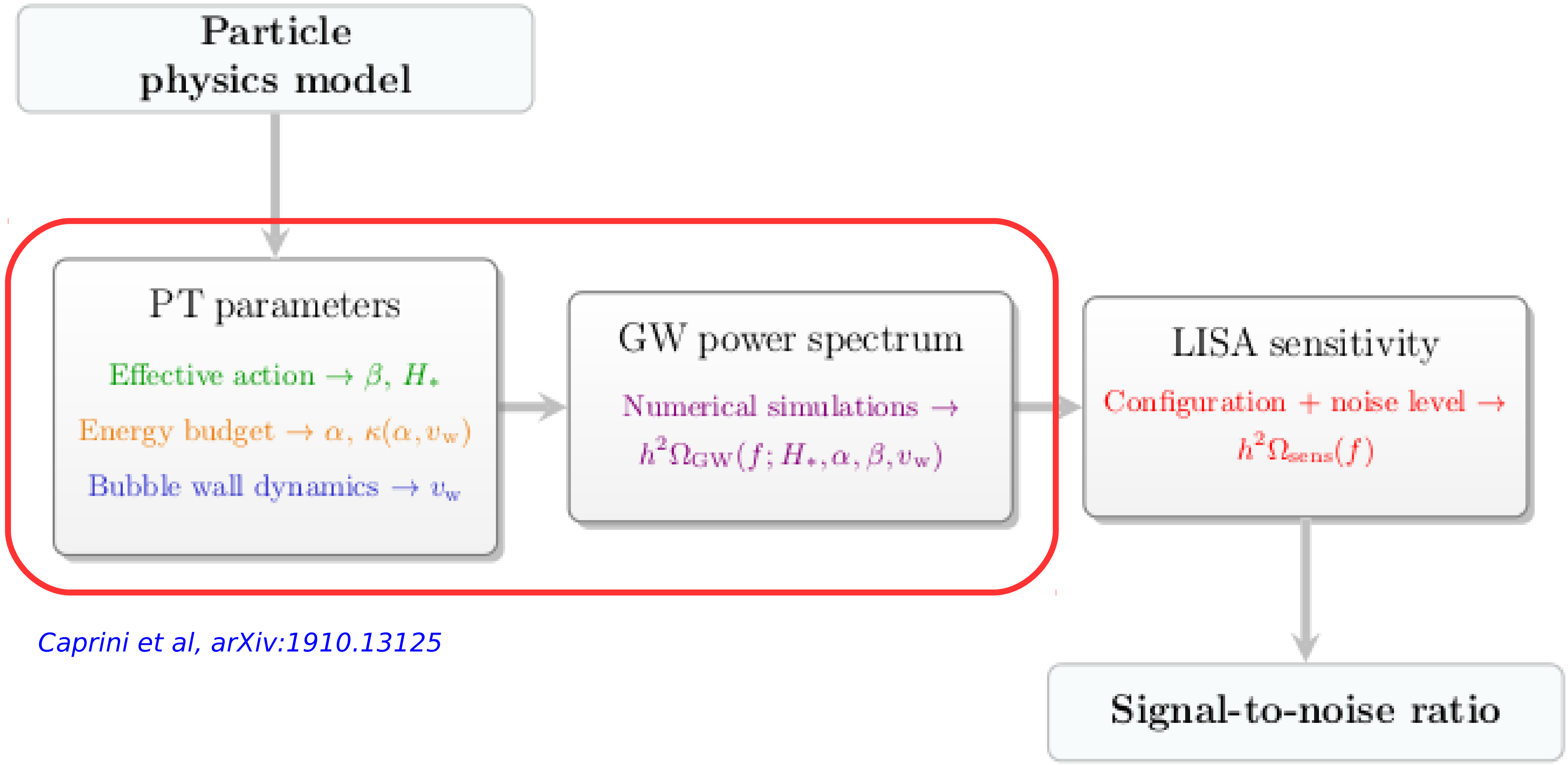
Configuration + noise level → $h^2\Omega_{sens}(f)$

Signal-to-noise ratio

Caprini et al, arXiv:1910.13125

Assess the capability of LISA to probe GW signal from EW epoch ⇨ **BSM physics**

GW from the EW Phase Transition with **LISA**



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Assess the capability of LISA to probe GW signal from EW epoch ⇨ BSM physics

Need to predict GW signal as robustly as possible

Thermal EW Phase Transition

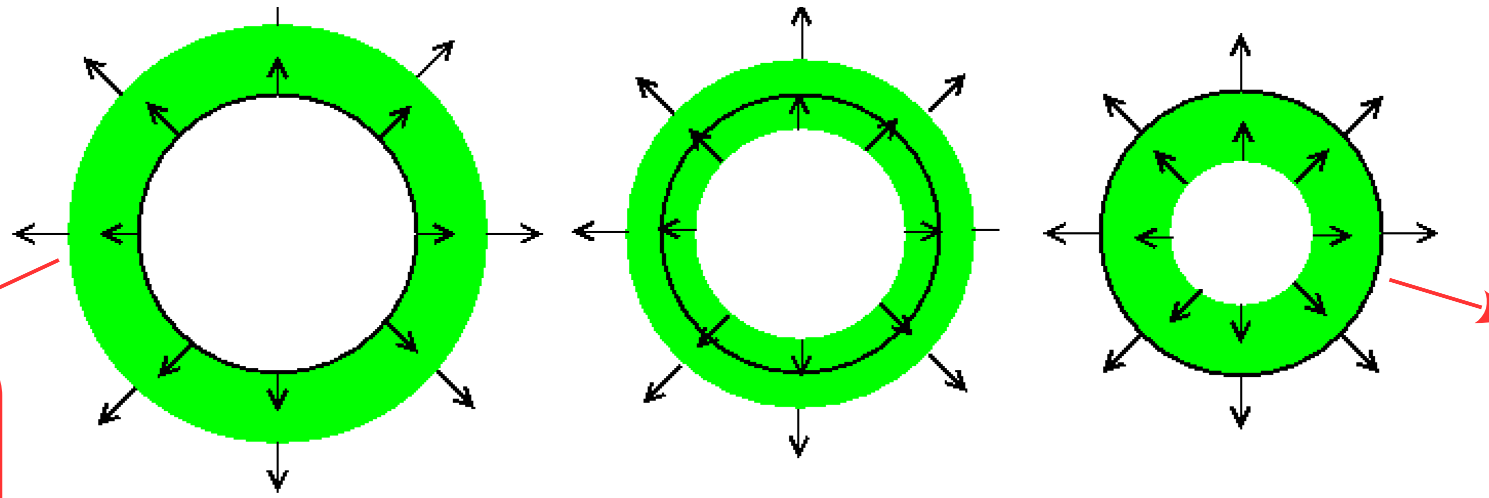
Energy liberated from phase change **transferred** (mostly) **to plasma**

- Kinetic energy ⇨ Thermal plasma bulk motion
- Thermal energy ⇨ Thermal plasma gets heated up

Depending on Higgs bubble wall velocity, energy transfer to plasma creates different types of **expanding fluid shells**

*Laine, Phys. Rev. D**49** (1994) 3847*

*Espinosa, Konstandin, No, Servant, JCAP **1006** (2010) 028*



$\mathbf{v}_W < \mathbf{c}_S$

fluid in front of phase boundary gets pushed

DEFLAGRATION

$\mathbf{v}_J > \mathbf{v}_W > \mathbf{c}_S$

HYBRID

$\mathbf{v}_W > \mathbf{v}_J > \mathbf{c}_S$

fluid behind phase boundary gets dragged

DETONATION

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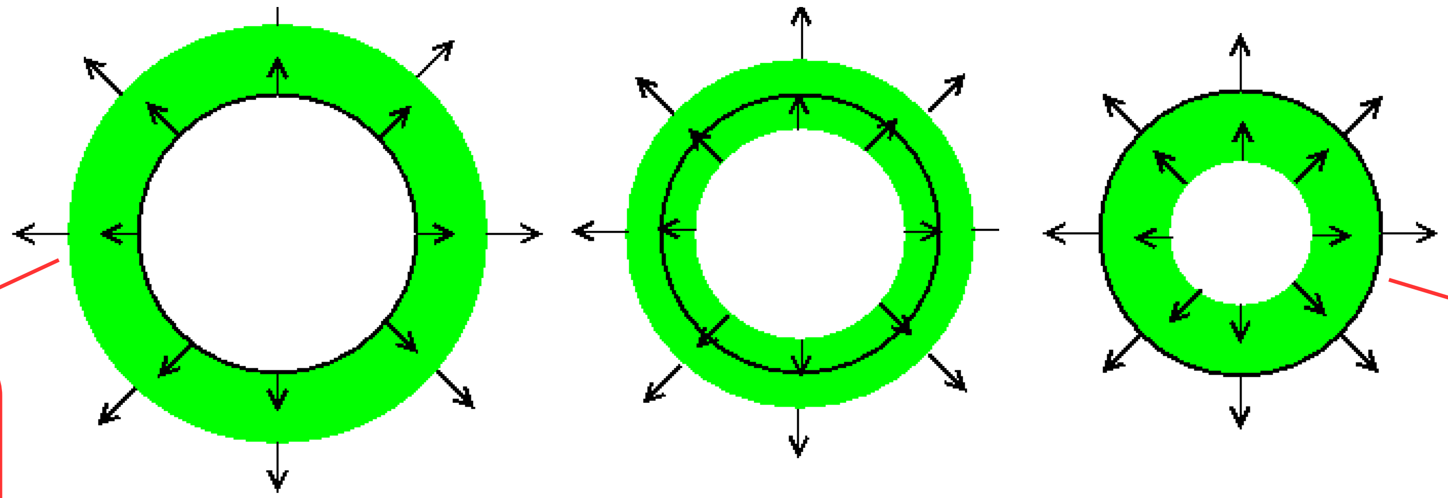

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DEFLAGRATION

$v_J > v_W > c_S$

HYBRID

$v_W > v_J > c_S$

fluid behind phase boundary gets dragged

DETONATION

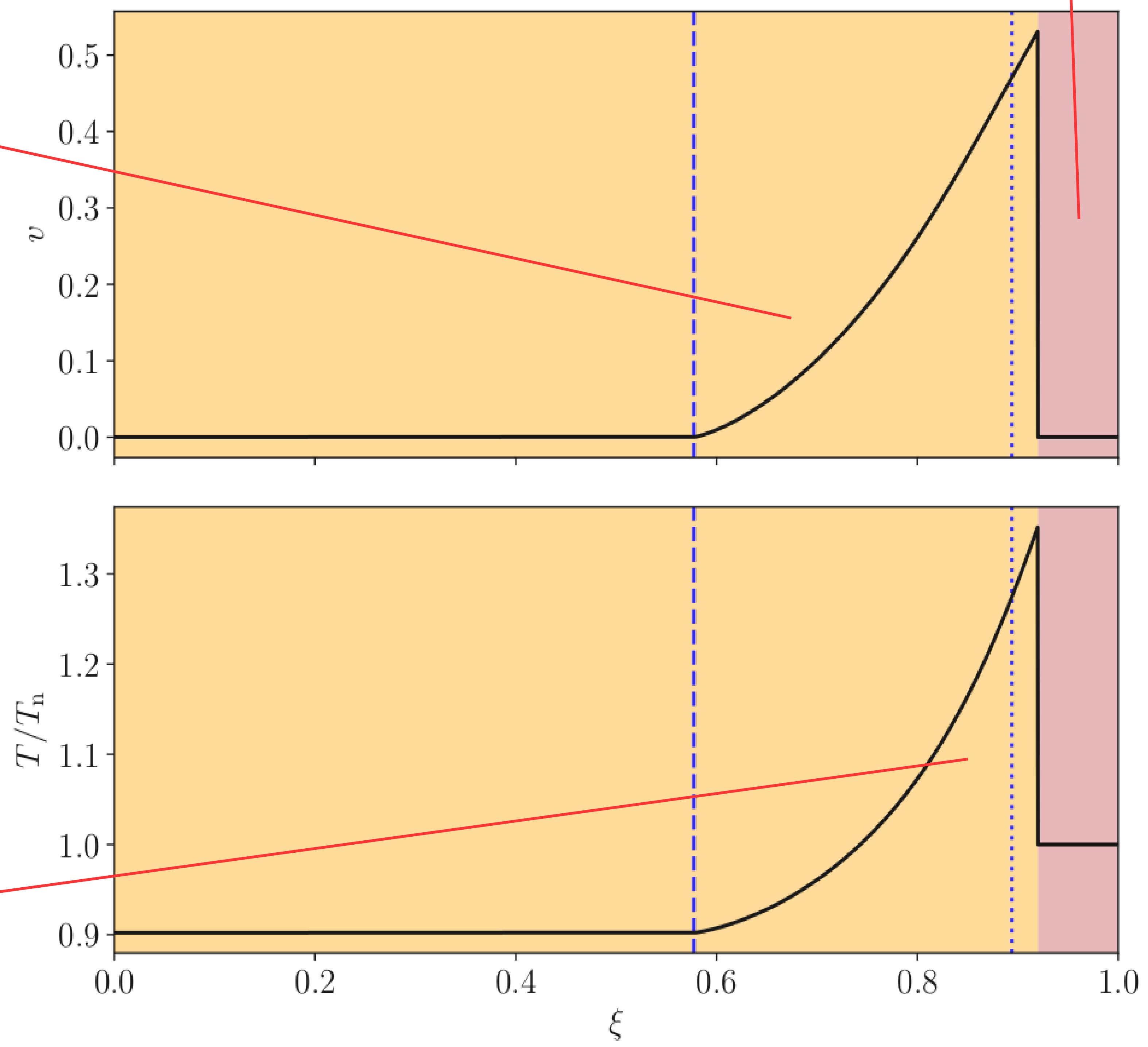
“Self-similar”

Fluid velocity and temperature depend on $\xi = r/t$

Detonations ($v_w > v_J > c_s$)

fluid at rest in front
of bubble

fluid behind phase boundary
gets dragged, until it becomes
at rest again



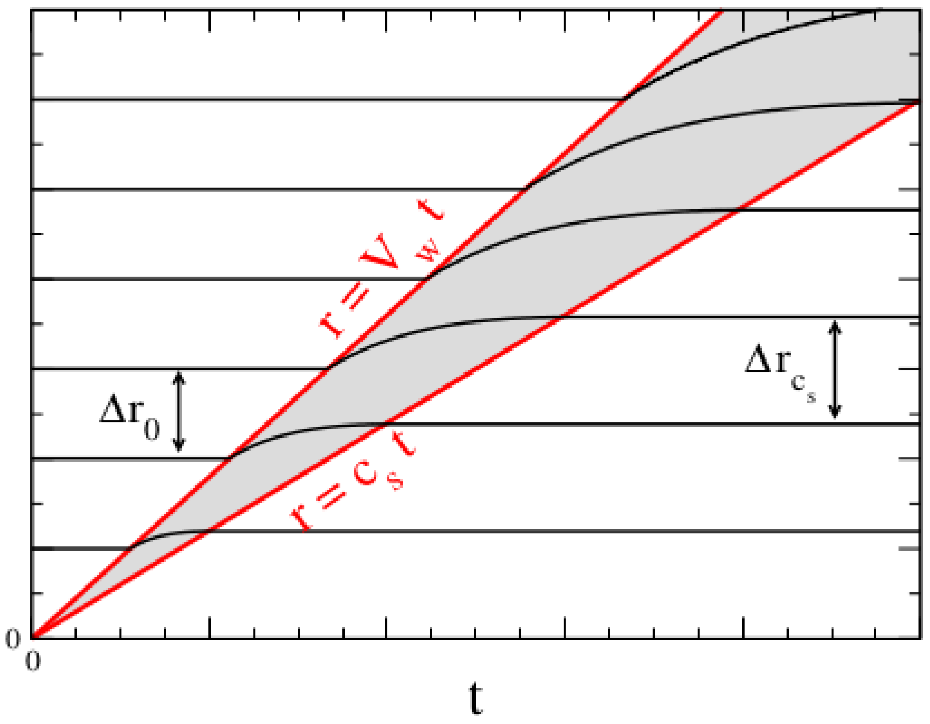
fluid heated inside bubble

Courtesy of D. Cutting (Sussex)

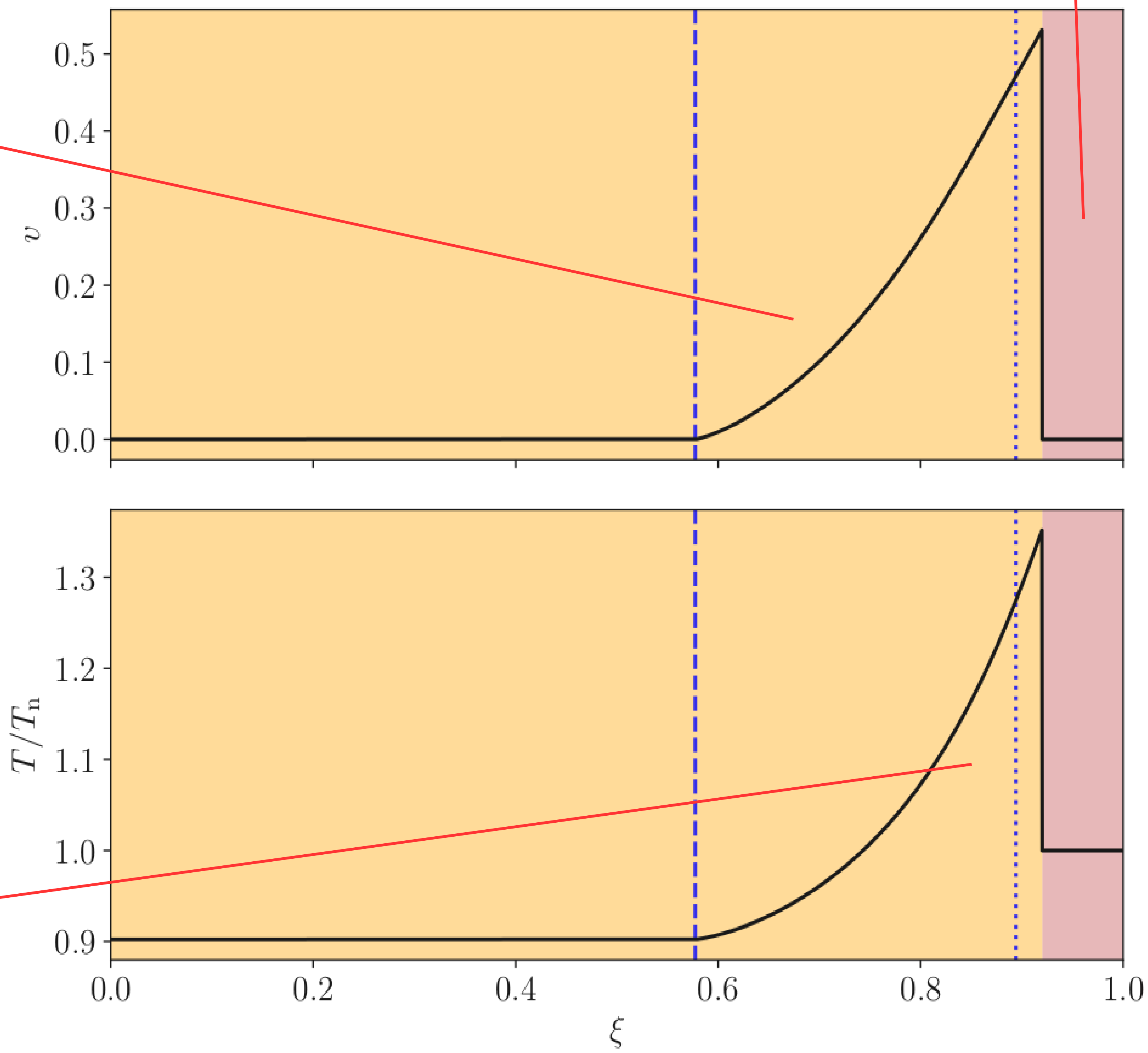
Detonations ($v_w > v_J > c_s$)

fluid at rest in front of bubble

fluid behind phase boundary gets dragged, until it becomes at rest again



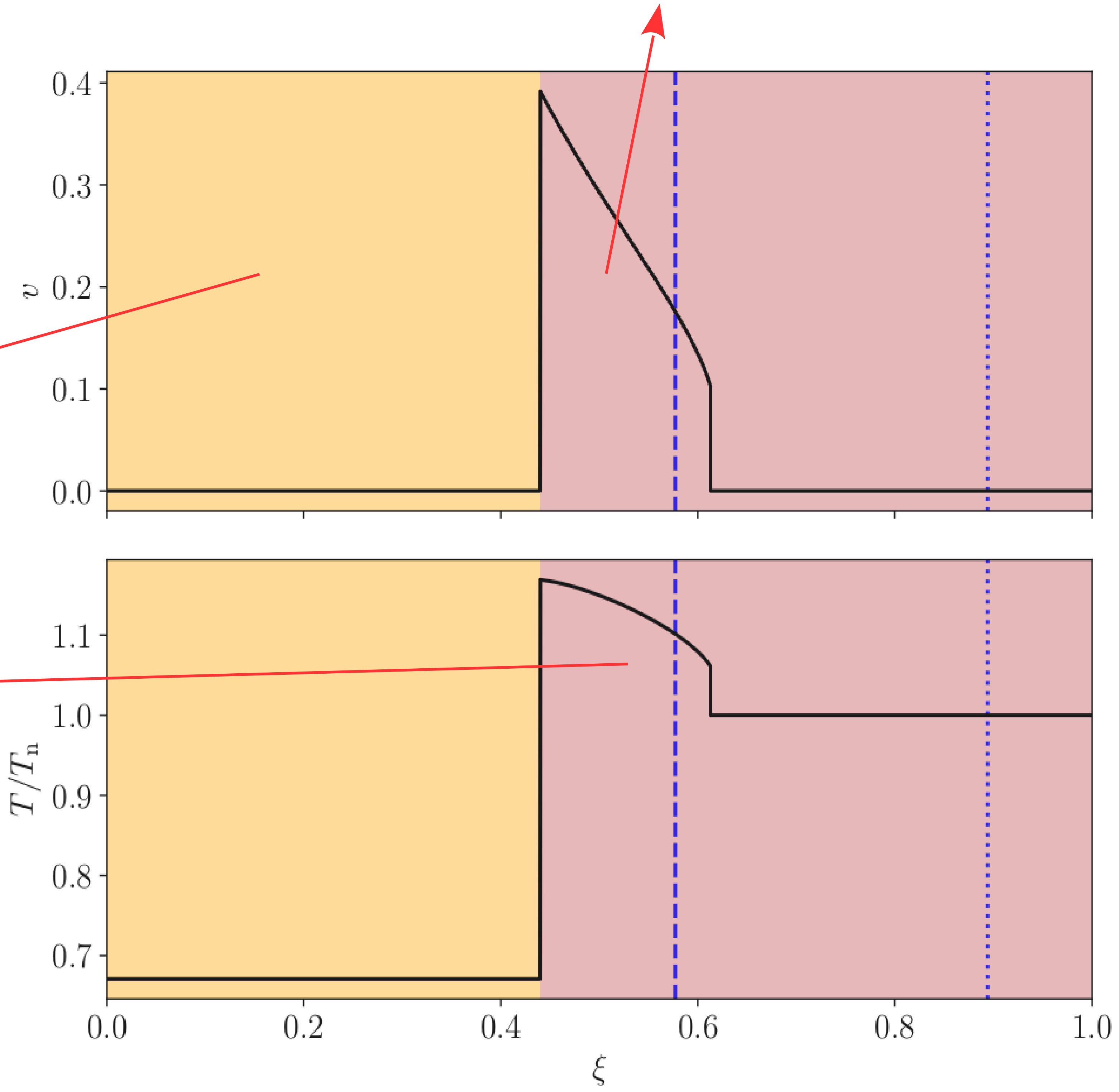
Caprini, No, JCAP ***1201*** *(2012) 031*



fluid heated inside bubble

Courtesy of D. Cutting (Sussex)

Deflagrations ($v_w < c_s$)



fluid in front of bubble
gets pushed

fluid at rest inside bubble

fluid heated in front
of bubble

Courtesy of D. Cutting (Sussex)

Fluid shell Profiles

$$\partial^\mu T_{\mu\nu}^{\text{plasma}} = 0$$

(with appropriate boundary conditions on bubble wall)

Local Thermal Equilibrium

$$T_{\mu\nu}^{plasma} = w\, u_\mu u_\nu - g_{\mu\nu}\, p$$

$$w = e + p$$

$$u_\mu = \frac{(1, \mathbf{v})}{\sqrt{1 - \mathbf{v}^2}} = (\gamma, \gamma\mathbf{v})$$

Self-similarity

$$v(r,t) = v(\xi = r/t)$$

Estimate of Energy available for GW production

(fluid bulk motion for one bubble)

$$\overline{U}_{\mathrm{f}}^{2} = \frac{3}{e\, v_{\mathrm{w}}^{3}} \int w(\xi)\, v^{2}\gamma^{2}\xi^{2} d\xi = \frac{\kappa\alpha}{1+\alpha}$$

(enthalpy weighted) plasma
RMS four velocity

*Hindmarsh, Huber, Rummukainen, Weir, PRD **96** (2017) 103520*

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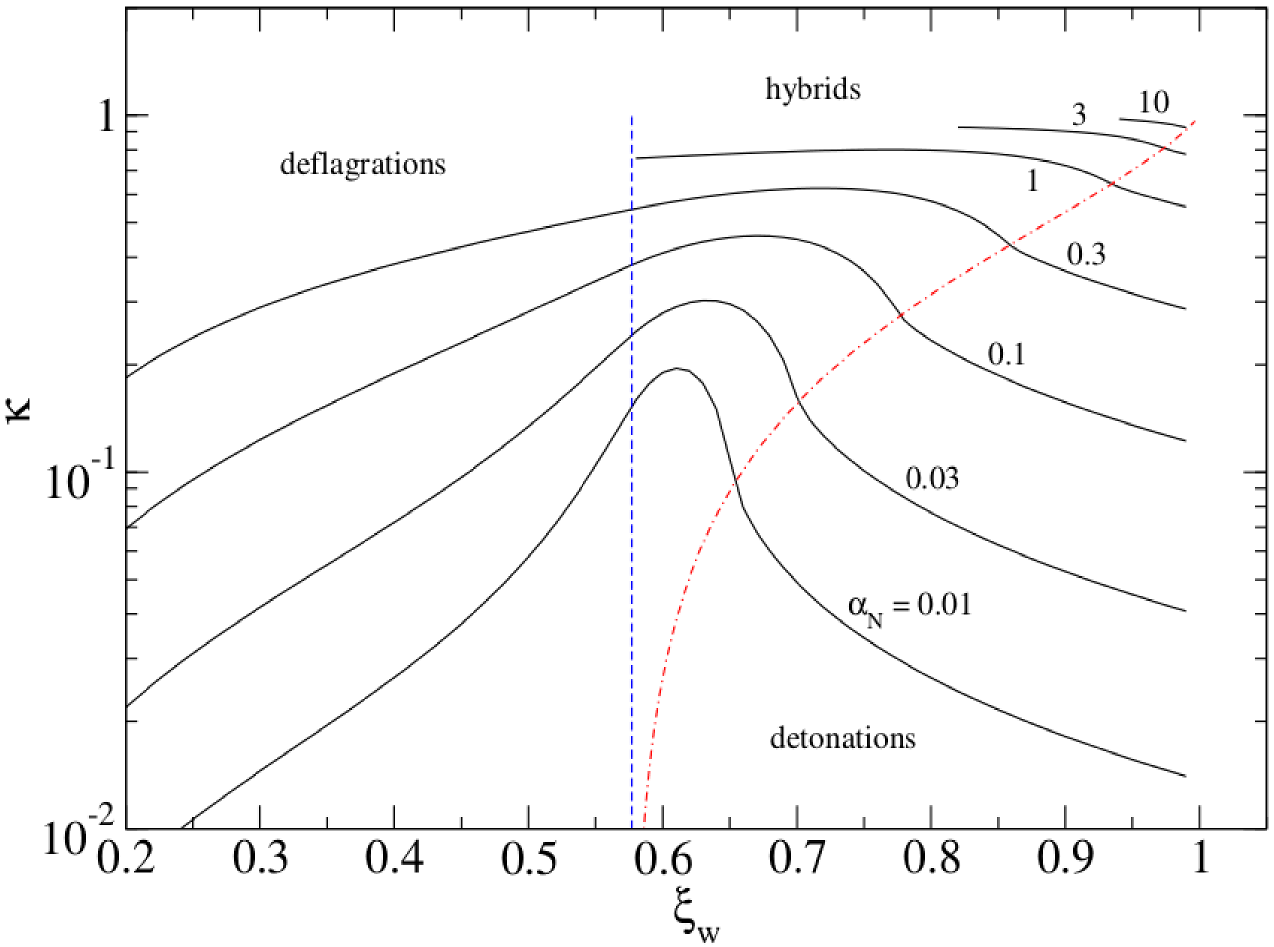
*Hindmarsh, Huber, Rummukainen, Weir, PRD **96** (2017) 103520*



Efficiency coefficient
(PT Energy Budget)

*Kamionkowski, Kosowsky, Turner, PRD **49** (1994) 2837*

*Espinosa, Konstandin, No, Servant, JCAP **1006** (2010) 028*

Gravitational Waves from Phase Transitions

- ❑ Gravitational waves (GWs) produced by several sources in a PT:

$$h^2\Omega_{\rm gw} = h^2\Omega_{\phi} + h^2\Omega_{\rm sw} + h^2\Omega_{\rm turb}$$

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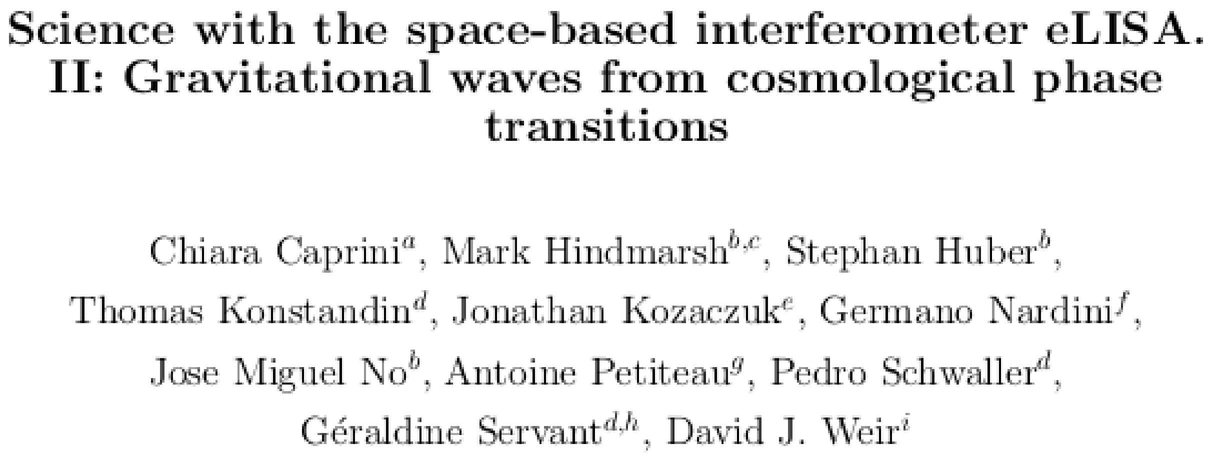
LISA Cosmology Working Group effort to provide state-of-art:

2015 CosWG Review

Caprini et al, JCAP ***1604*** *(2016) 001*

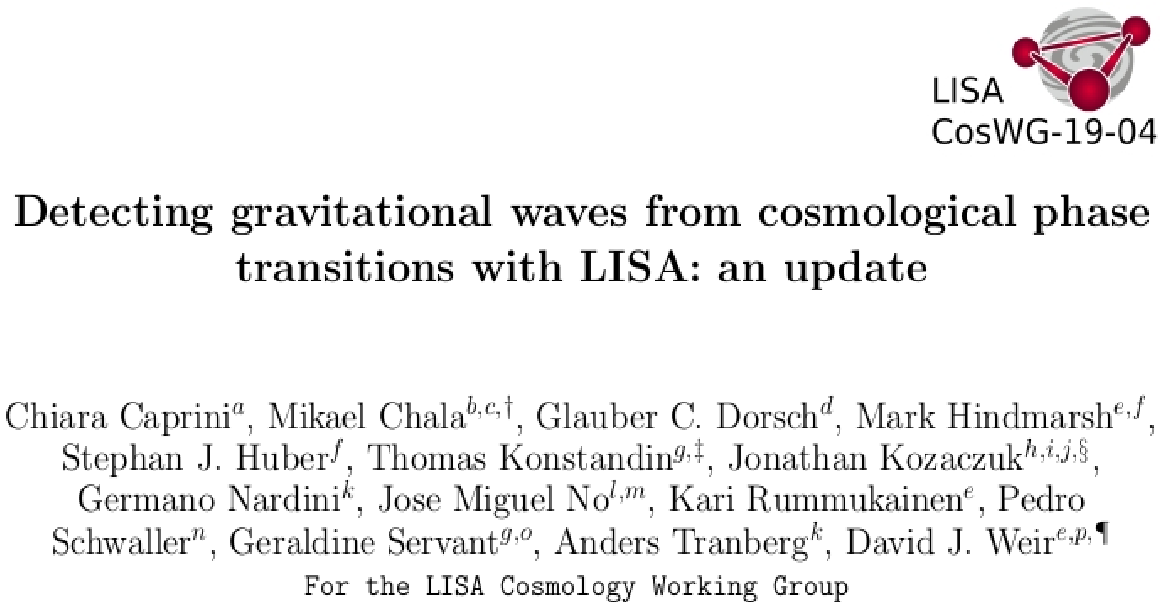


Science with the space-based interferometer eLISA. II: Gravitational waves from cosmological phase transitions

Chiara Caprini[a], Mark Hindmarsh[b,c], Stephan Huber[b],
Thomas Konstandin[d], Jonathan Kozaczuk[e], Germano Nardini[f],
Jose Miguel No[b], Antoine Petiteau[g], Pedro Schwaller[d],
Géraldine Servant[d,h], David J. Weir[i]

+ very recent update

Caprini et al, arXiv:1910.13125



LISA
CosWG-19-04

Detecting gravitational waves from cosmological phase transitions with LISA: an update

Chiara Caprini[a], Mikael Chala[b,c,†], Glauber C. Dorsch[d], Mark Hindmarsh[e,f],
Stephan J. Huber[f], Thomas Konstandin[g,‡], Jonathan Kozaczuk[h,i,j,§],
Germano Nardini[k], Jose Miguel No[l,m], Kari Rummukainen[e], Pedro
Schwaller[n], Geraldine Servant[g,o], Anders Tranberg[k], David J. Weir[e,p,¶]
For the LISA Cosmology Working Group

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*Kosowsky, Turner, Watkins, PRL **69** (1992) 2026; PRD **45** (1992) 4514*
*Huber, Konstandin, JCAP **0809** (2008) 022*
*Weir, PRD **93** (2016) 124037*
*Cutting, Hindmarsh, Weir, PRD **97** (2018) 123513*

In general, negligible expect for very strong supercooling ⇒ $\alpha >> 1$

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*Cutting, Hindmarsh, Weir, PRD **97** (2018) 123513*

In general, negligible expect for very strong supercooling ⇒ α >> 1

Such amount of supercooling incompatible with PT completion...
*Ellis, Lewicki, No, JCAP **1904** (2019) 003*

...except for conformal scalar potentials!

Iason Baldes talk coming up!

Gravitational Waves from Phase Transitions

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Hindmarsh, Huber, Rummukainen, Weir, PRL ***112*** *(2014) 041301; PRD* ***92*** *(2015) 123009; PRD* ***96*** *(2017) 103520*
Hindmarsh, PRL ***120*** *(2018) 071301*
Konstandin, JCAP ***1803*** *(2018) 047*
Hindmarsh, Hijazi, arXiv:1909.10040

Typically dominant signal

GW power spectrum (numerical simulations)

$$\boxed{\frac{d\Omega_{\rm gw,0}}{d\ln(f)} = \boxed{0.687 F_{\rm gw,0} K^2 (H_* R_*/c_s)\tilde{\Omega}_{\rm gw}}\,\boxed{C\left(\frac{f}{f_{\rm p,0}}\right)}}$$

Peak amplitude
(at maximum)

Peak frequency
(at maximum)

Spectral Shape
(Broken Power Law)

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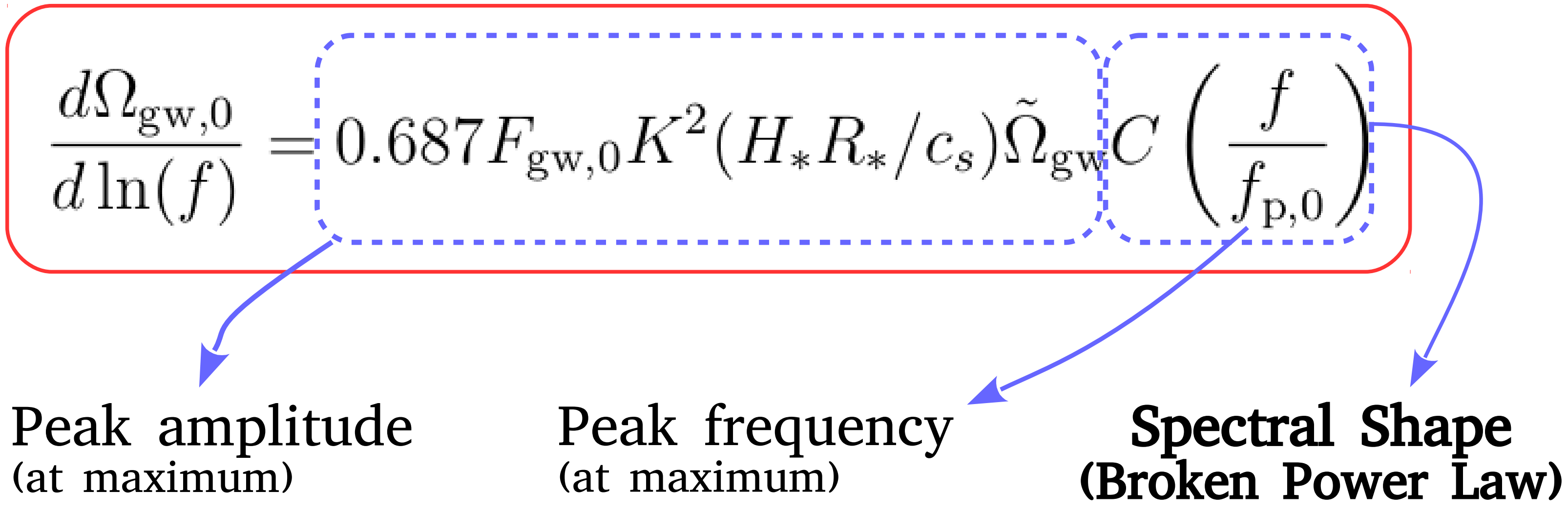

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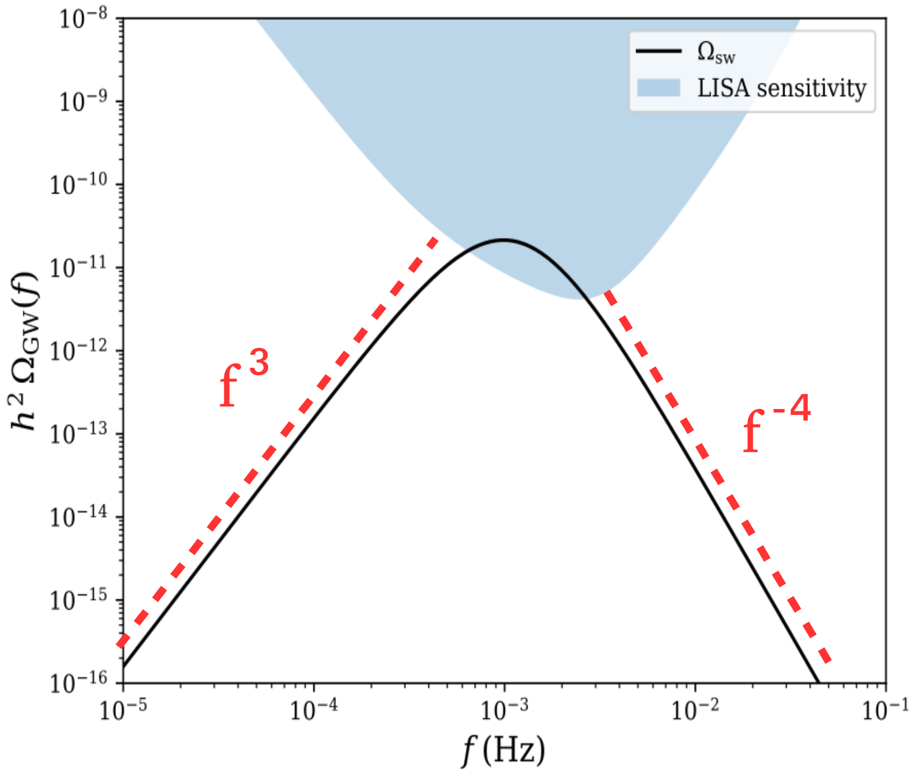
Peak amplitude
(at maximum)

Peak frequency
(at maximum)

Spectral Shape
(Broken Power Law)

$$C(s) = s^3\left(\frac{7}{4+3s^2}\right)^{\frac{7}{2}}$$



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$$F_{\rm gw,0} = (3.57 \pm 0.05) \times 10^{-5} \left(\frac{100}{g_*}\right)^{\frac{1}{3}}$$

$$\tilde{\Omega}_{\rm gw} \sim 10^{-2}$$

(from simulation)

$$K = \frac{\langle w\gamma^2 v^2\rangle}{\bar{e}} = \Gamma \bar{U}_{\rm f}^2$$

Kinetic Energy Fraction

Gravitational Waves from Phase Transitions

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After $\tau_{\rm sh} \sim L_{\rm f}/\overline{U}_{\rm f}$, fluid becomes nonlinear (shock formation)

⇩

characteristic fluid length scale

Sound wave GW source shuts-off

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$H_*\tau_{\rm sh} = H_* R_*/K^{1/2} <$ 1

Gravitational Waves from Phase Transitions

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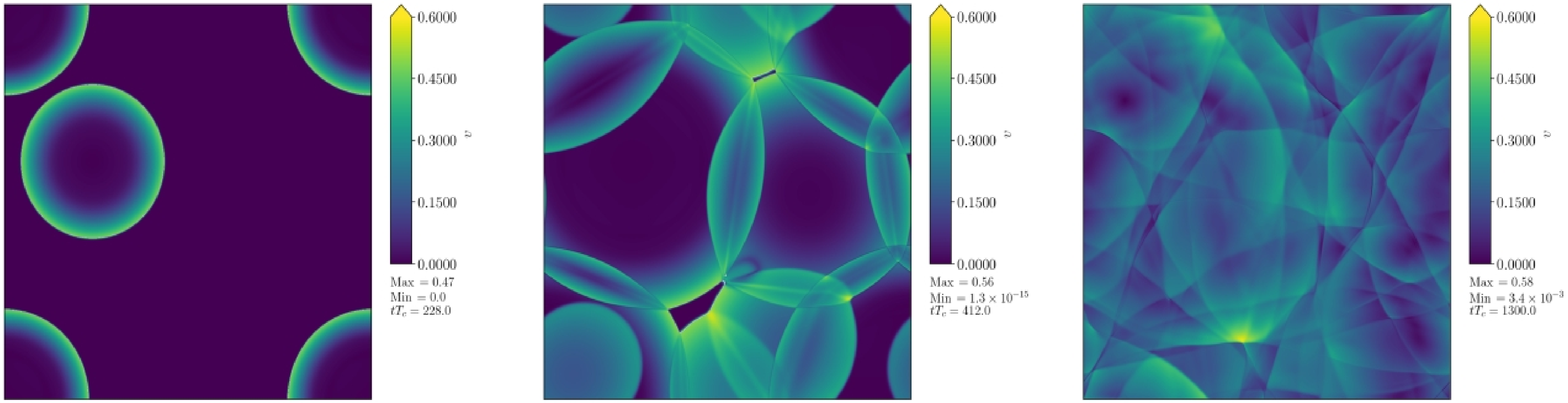

Cutting, Hindmarsh, Weir, arXiv:1906.00480

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⇩

Sound wave GW source shuts-off

- ☐ $h^2\Omega_{\rm turb}$ sourced by plasma turbulence (vortical modes)

*Gogoberidze, Kahniashvili, Kosowsky, PRD **76** (2007) 083002*
*Caprini, Durrer, Servant, JCAP **0912** (2009) 024*
Roper Pol, Mandal, Brandenburg, Kahniashvili, Kosowsky, arXiv:1903.08585

→ Turbulent flow expected to develop when sound waves shut-off

→ **Vorticity can also coexist with sound waves for deflagrations and $\alpha > 0.1$**

Cutting, Hindmarsh, Weir, arXiv:1906.00480

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Numerical simulations still ongoing

Gravitational Waves from Phase Transitions

Duration of sound wave GW source

Initially assumed linear fluid regime lasts approx. a Hubble time

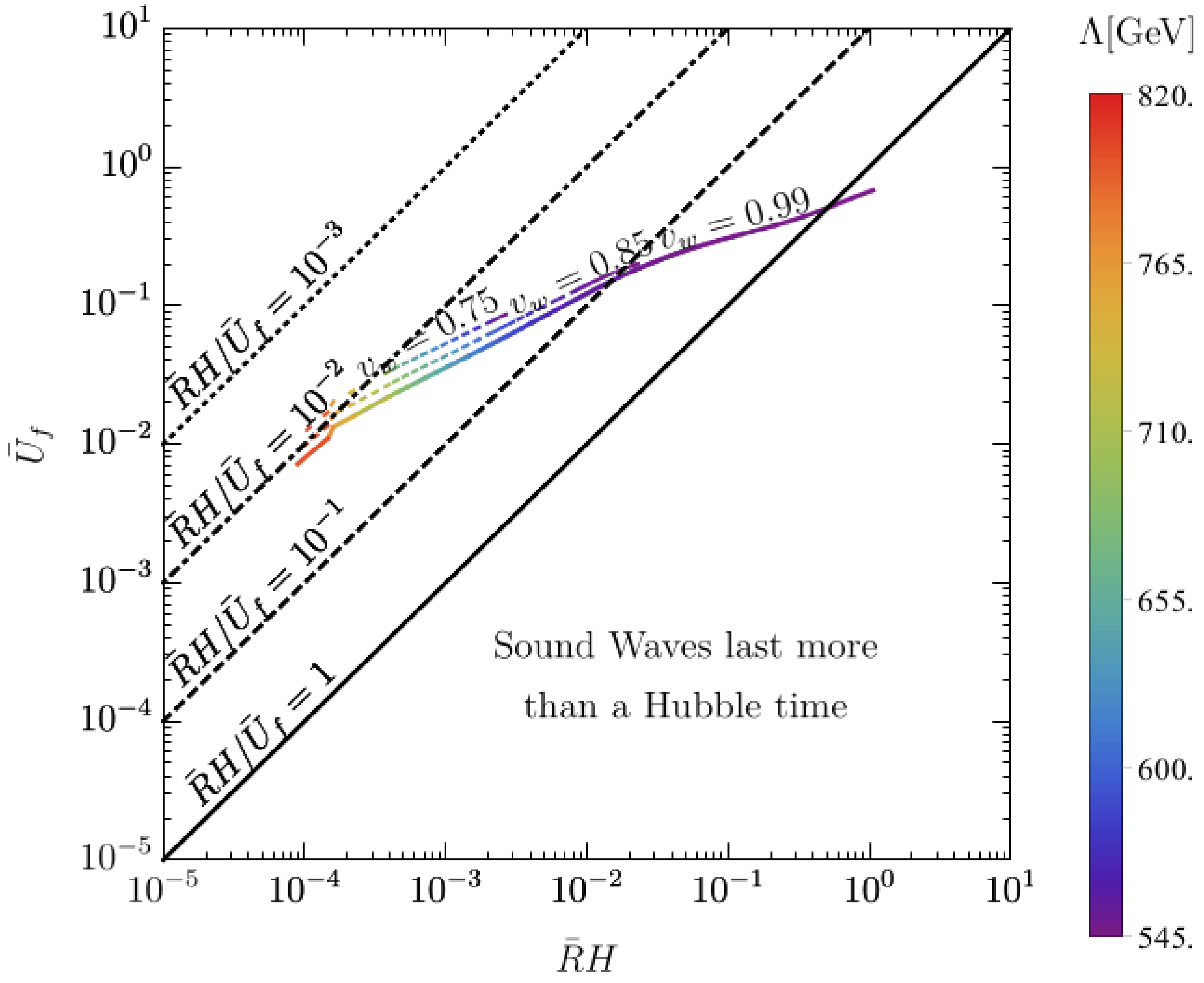
$$\tau_{\mathrm{sh}} \gtrsim H_*^{-1}$$

But non-linearities generally "cut short" the sound wave GW source:

*Ellis, Lewicki, No, JCAP **1904** (2019) 003*

Concrete BSM example:

$$V(H) = -m^2|H|^2 + \lambda|H|^4 + \frac{1}{\Lambda^2}|H|^6$$

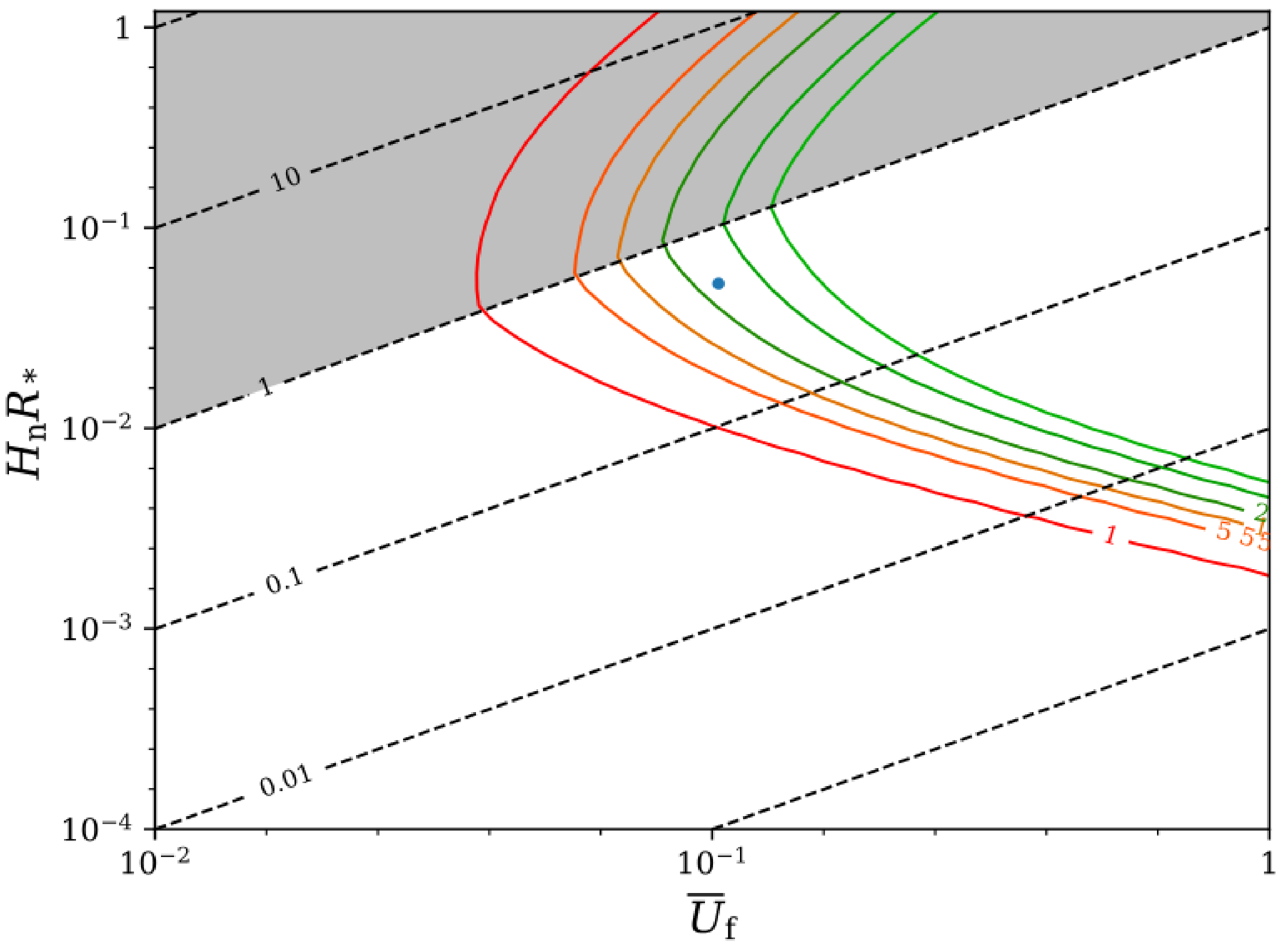


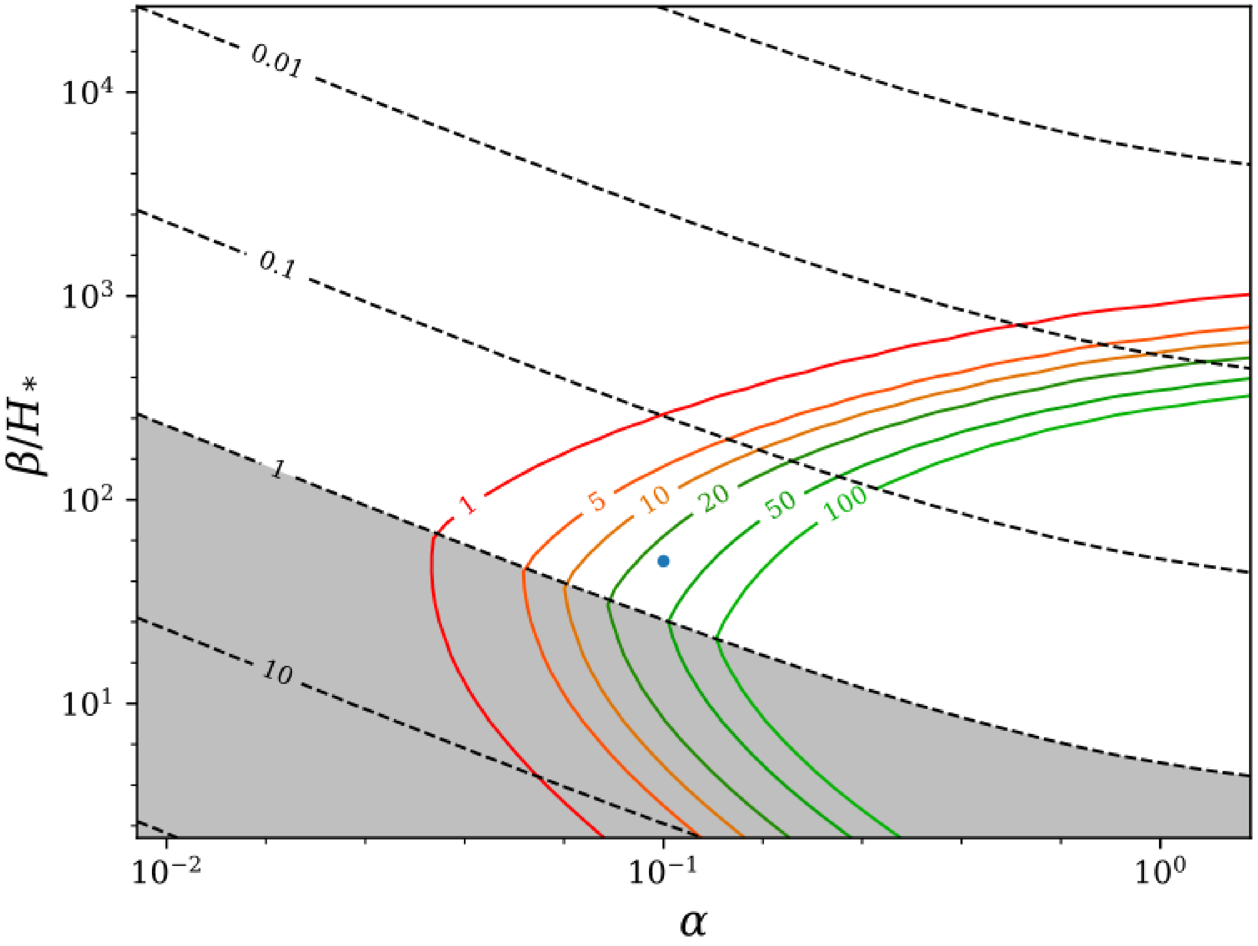
Gravitational Waves from Phase Transitions

LISA signal to noise

$$\mathrm{SNR} = \sqrt{\mathcal{T} \int_{f_{\mathrm{min}}}^{f_{\mathrm{max}}} \mathrm{d}f \left[\frac{h^2 \Omega_{\mathrm{GW}}(f)}{h^2 \Omega_{\mathrm{Sens}}(f)} \right]^2}$$

$$h^2 \Omega_{\mathrm{Sens}}(f) = \frac{2\pi^2}{3 H_0^2} f^3 S_h(f)$$





PTPlot Tool - D. Weir

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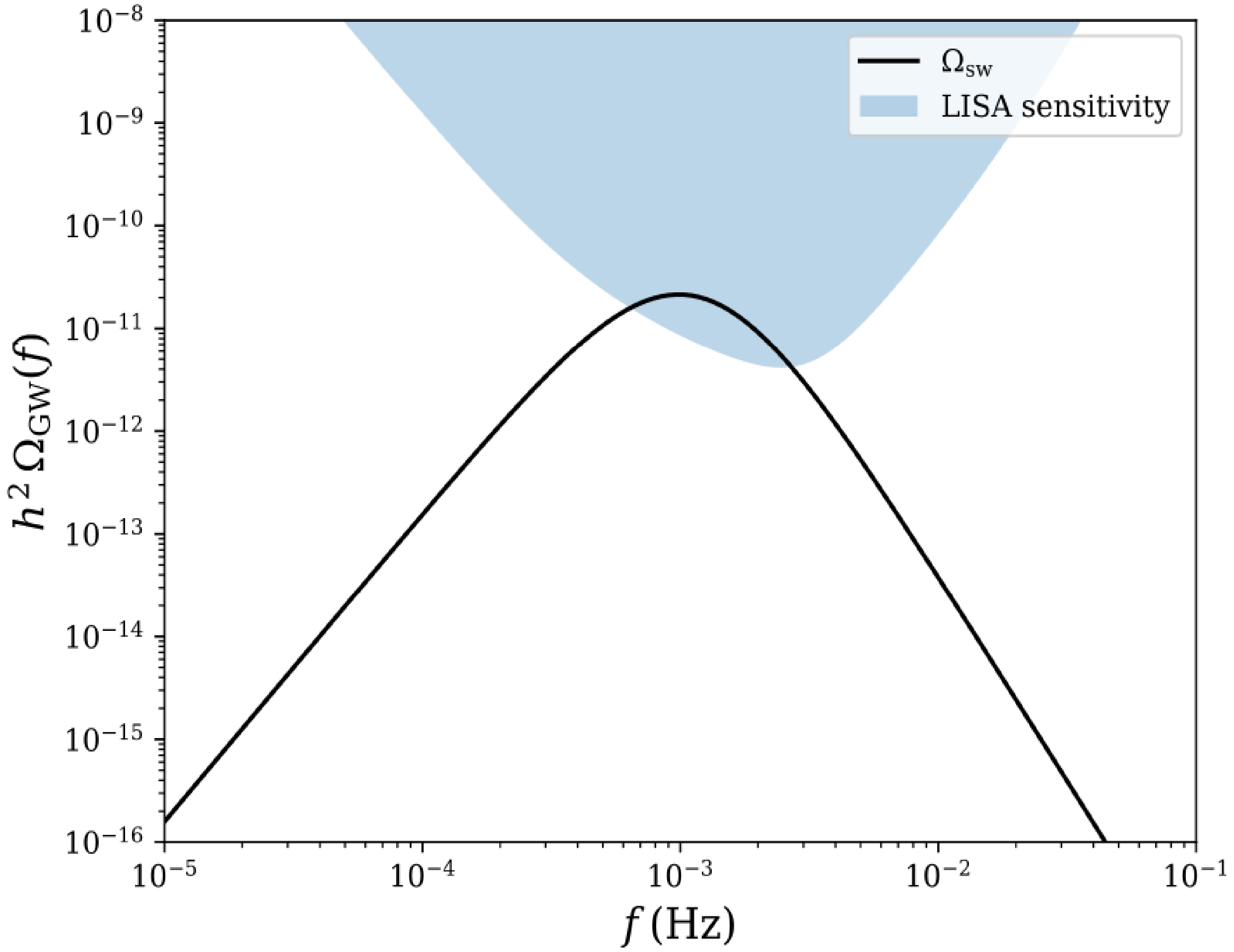
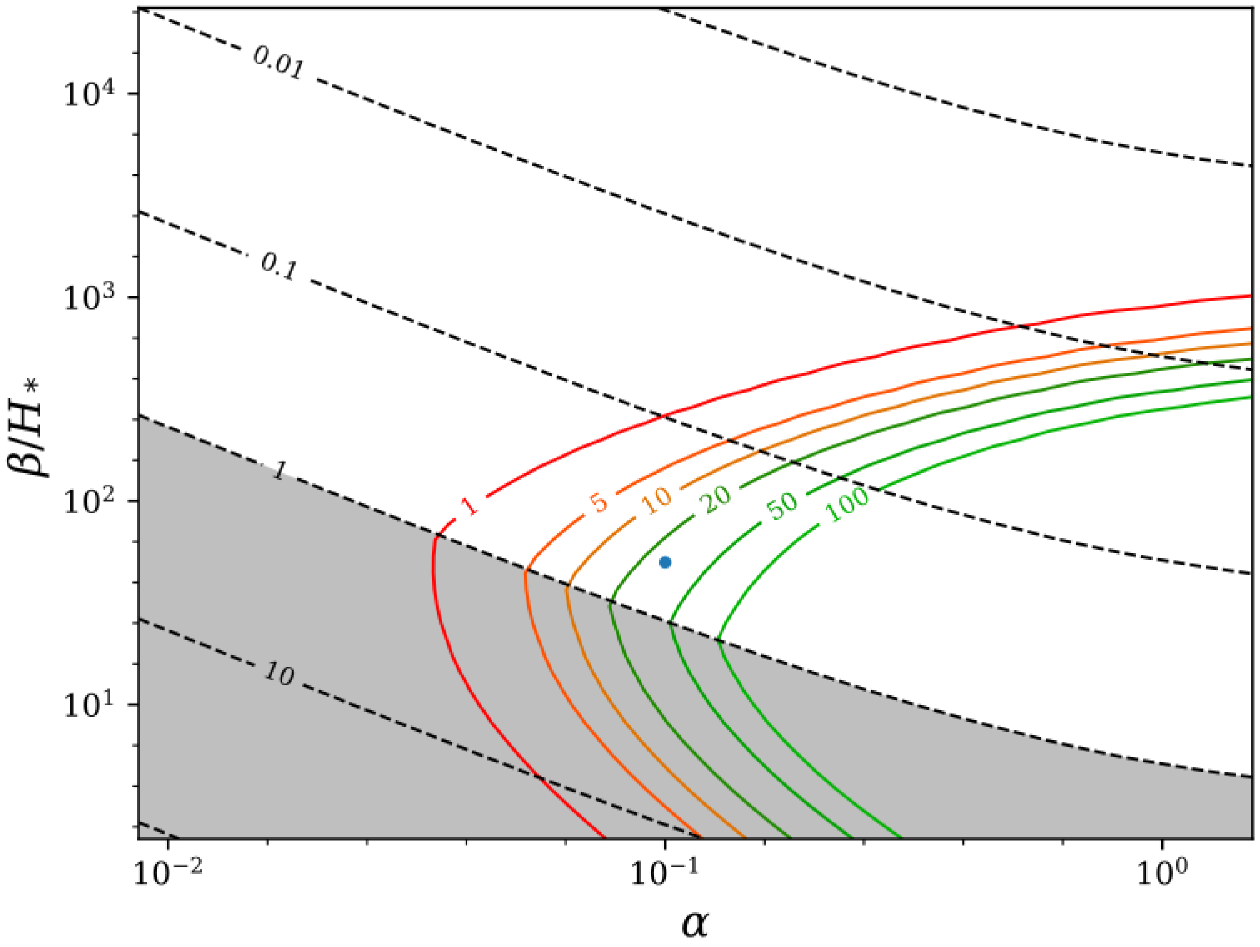




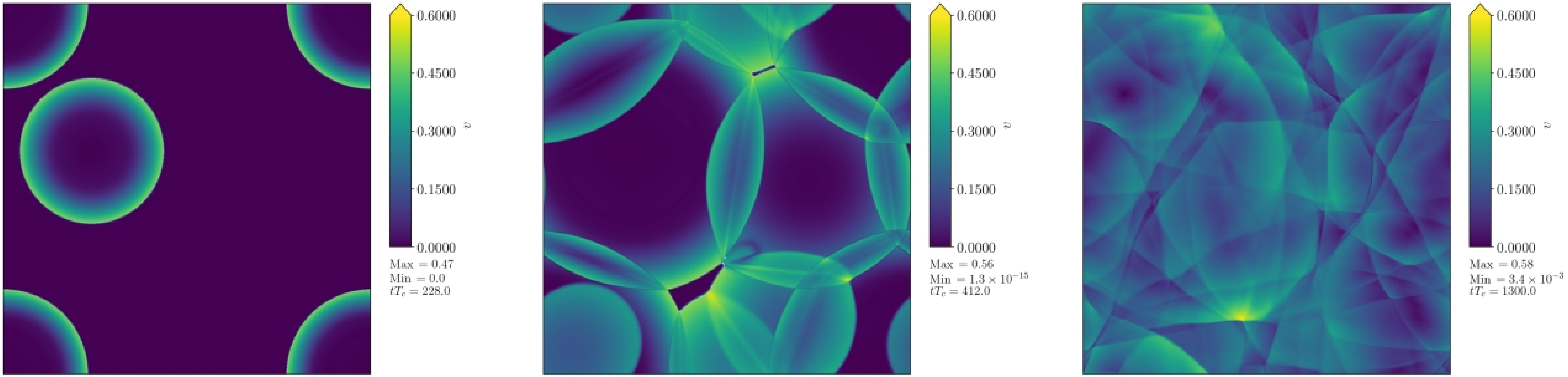
PTPlot Tool - D. Weir

Gravitational Waves from Phase Transitions

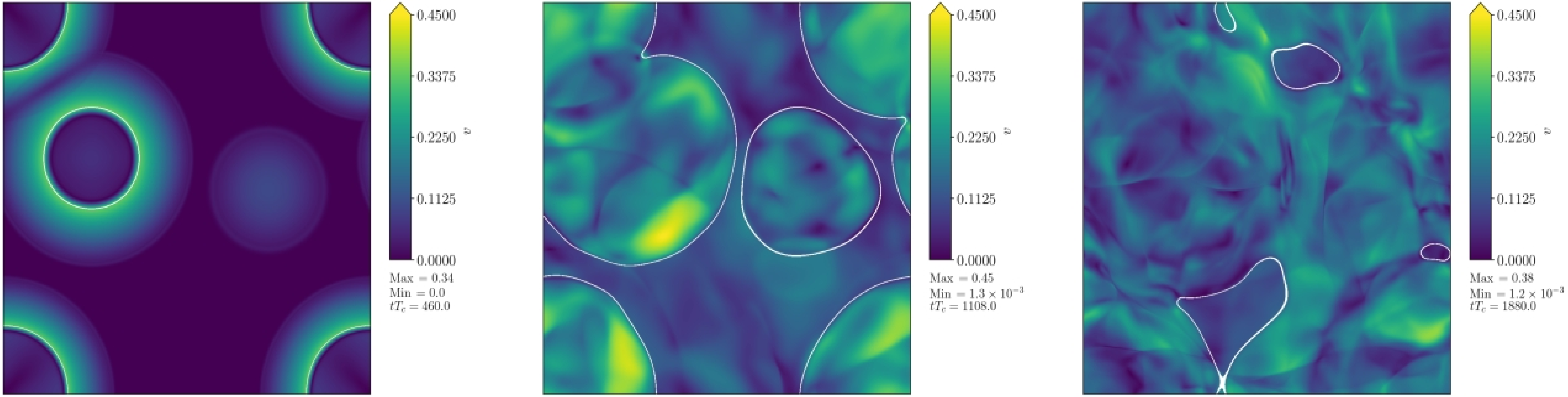
Understanding of vorticity generation is ongoing...

Cutting, Hindmarsh, Weir, arXiv:1906.00480

Detonations ($\alpha > 0.1$)



Deflagrations ($\alpha > 0.1$)



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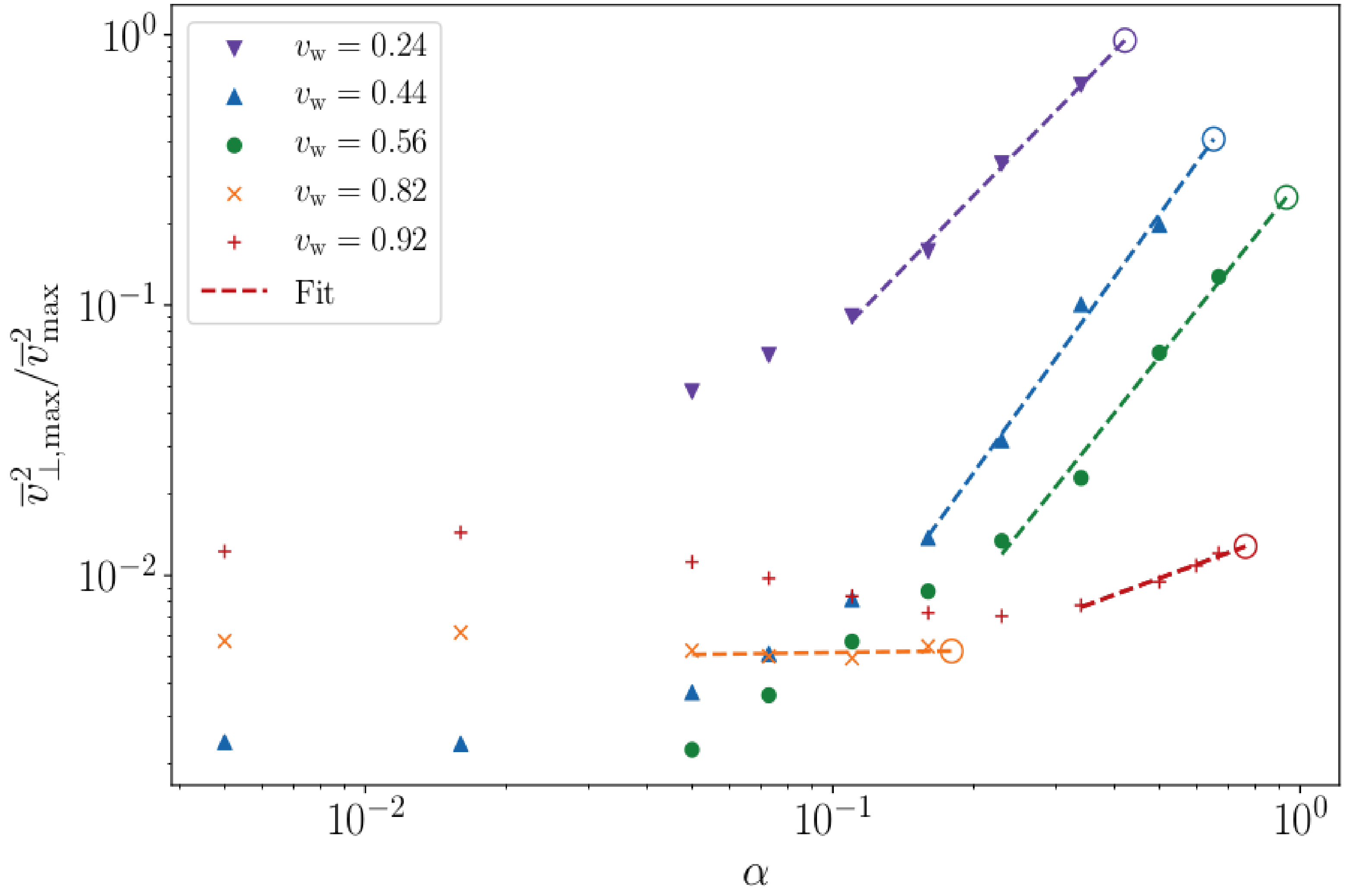


Deflagrations with large α (> 0.1) generate significant vorticity coexisting with sound waves!

In the last couple of min...

GW generation *vs* EW Baryogenesis in 1^{st} Order EW Phase Transition

In the last couple of min…

GW generation *vs* EW Baryogenesis in 1st Order EW Phase Transition

GWs: Sizable plasma bulk motion ⇨ **Sizable v_W**

EWBG: Velocities ~ 0.05 - 0.1 preferred

(efficient transport)

Incompatible?

In the last couple of min...

GW generation *vs* EW Baryogenesis in 1st Order EW Phase Transition

GWs: Sizable plasma bulk motion ⇨ **Sizable v_w**

EWBG: Velocities ~ 0.05 - 0.1 preferred

(efficient transport)

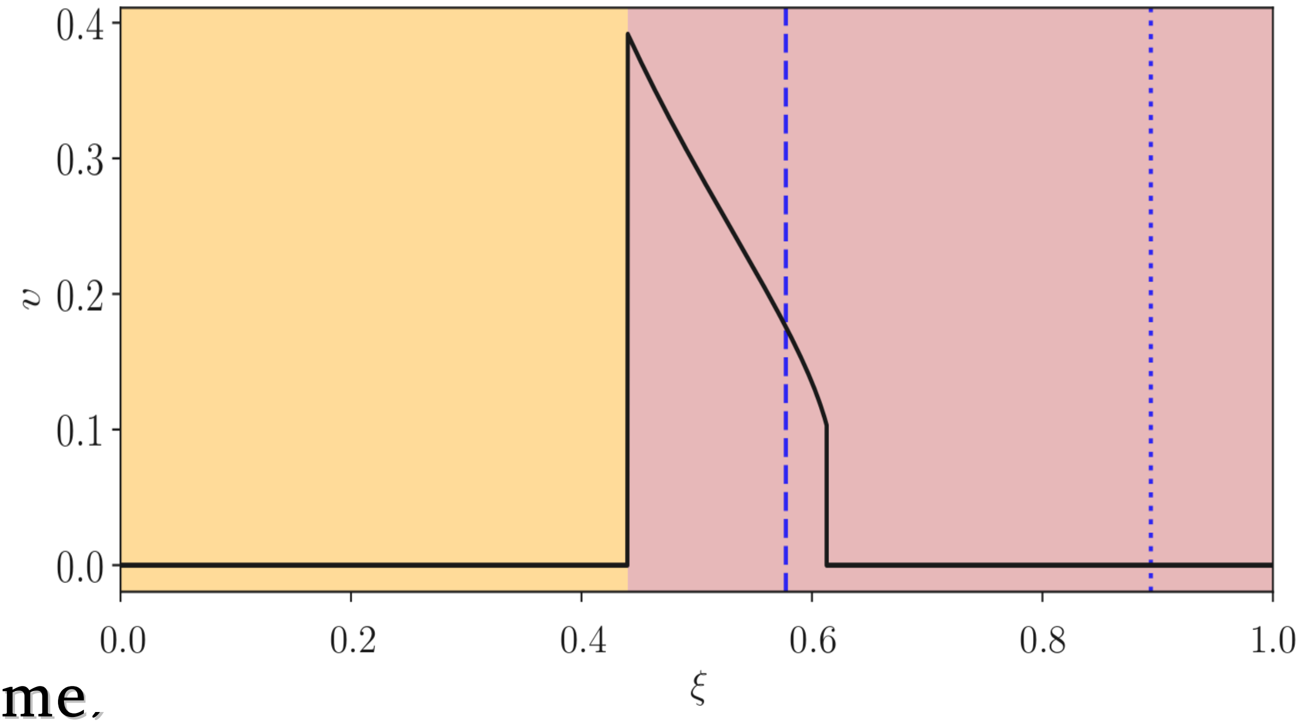
Incompatible?

NO!

Relevant velocities are not the same...

(for deflagrations)

*No, PRD **84** (2011) 124025*



GWs: Fluid velocity (bubble rest frame)

EWBG: Fluid velocity (bubble wall rest frame

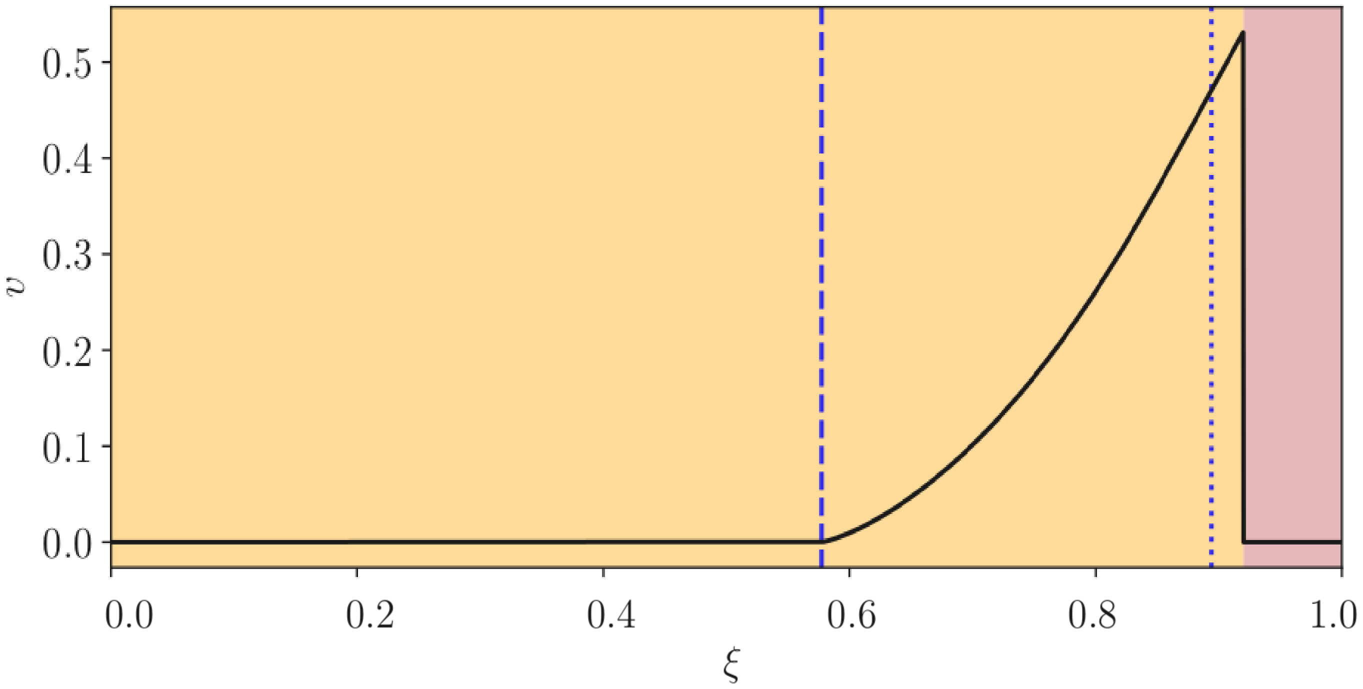
(relative velocity between bubble wall and plasma in front)

In the last couple of min...

GW generation *vs* EW Baryogenesis in 1^{st} Order EW Phase Transition

For detonations: EWBG would not work

(inefficient transport)



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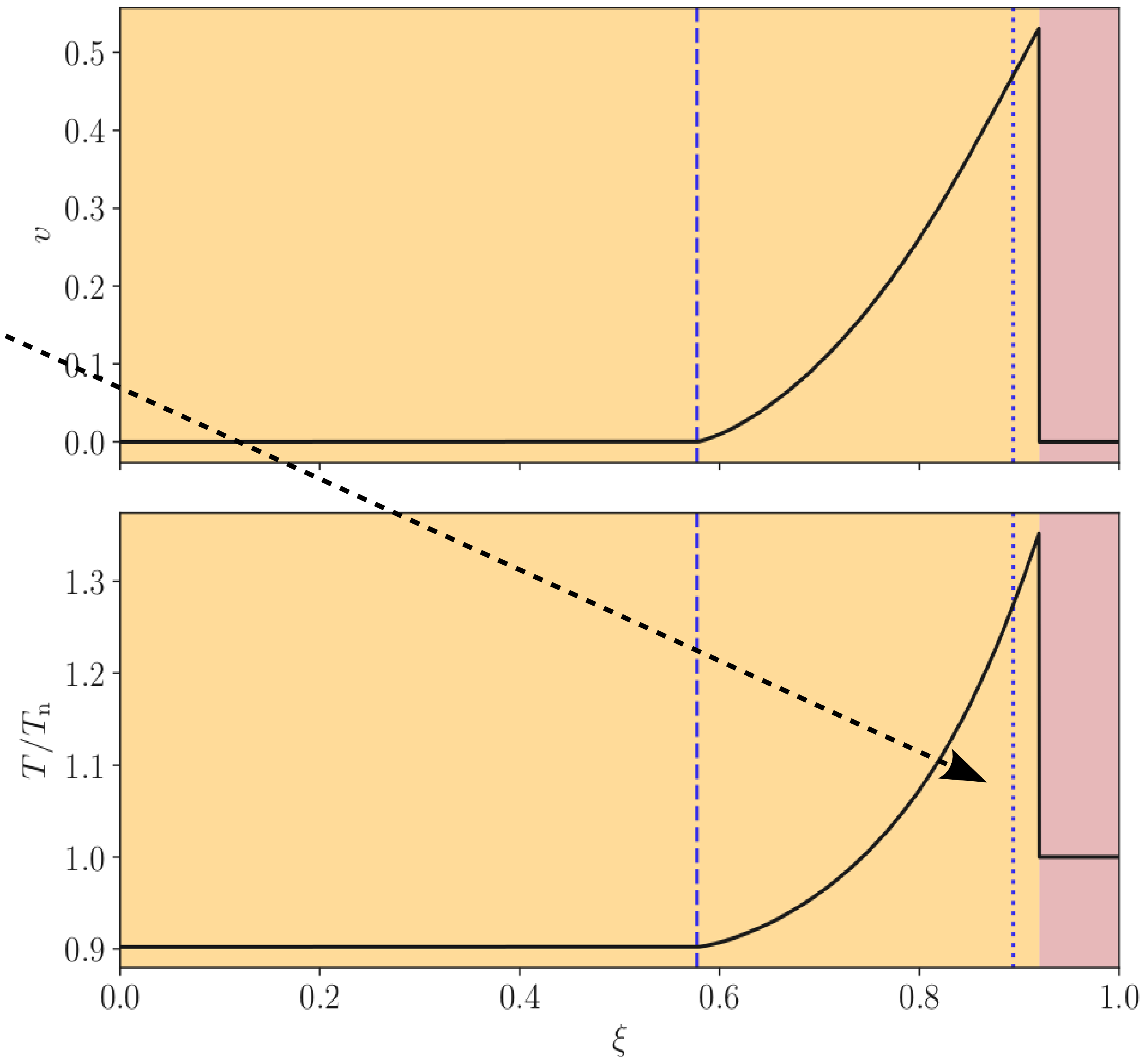
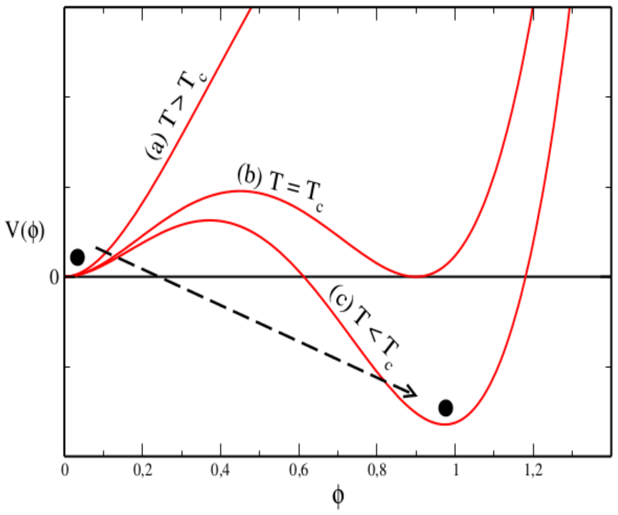
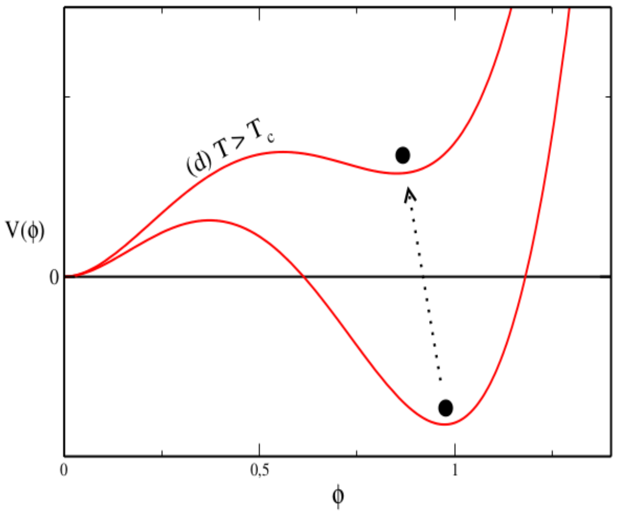


Possible to do EWBG from local back-tunneling (if T > T_C) in broken phase due to reheating!

*Caprini, No, JCAP **1201** (2011) 031*





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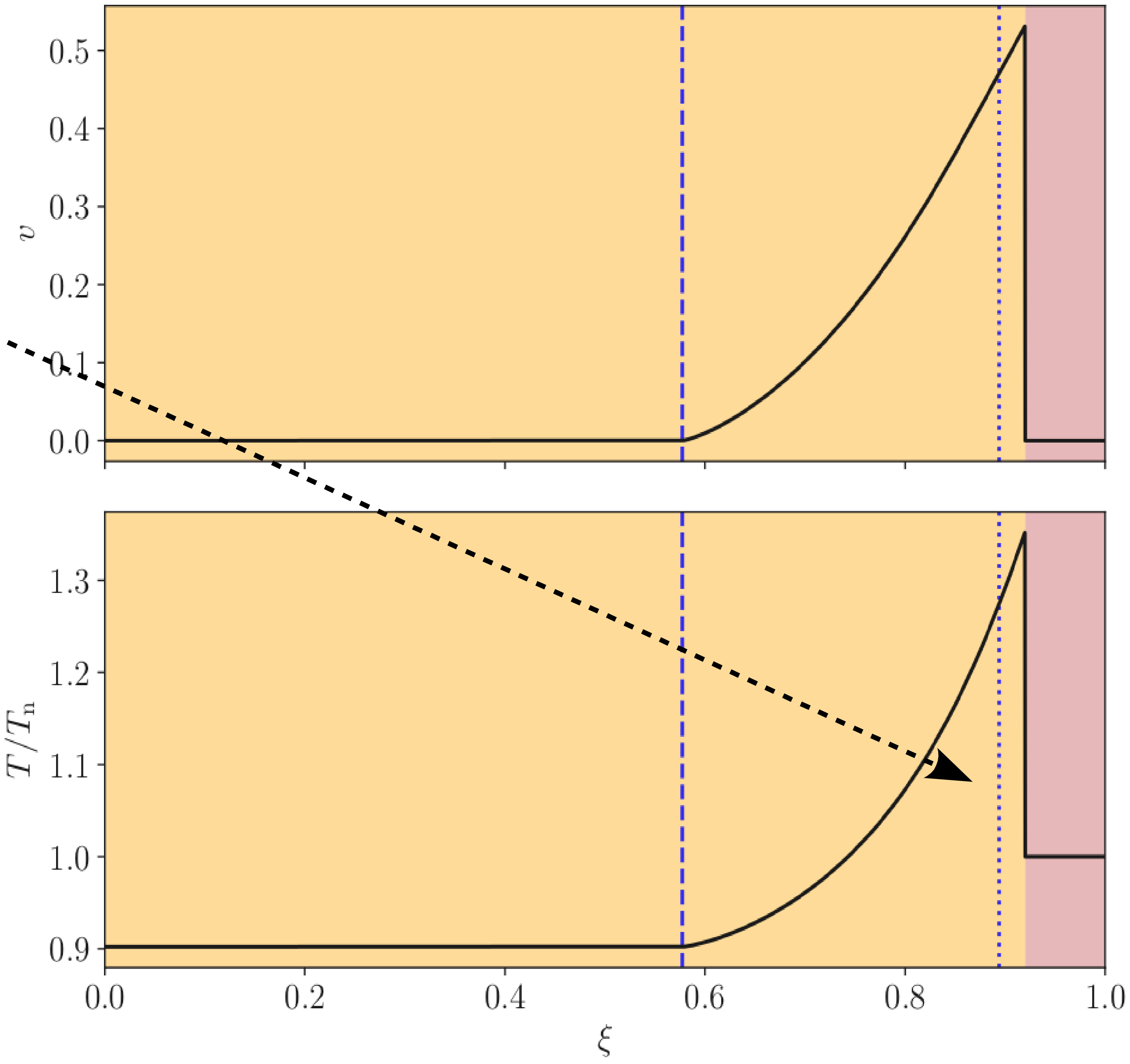

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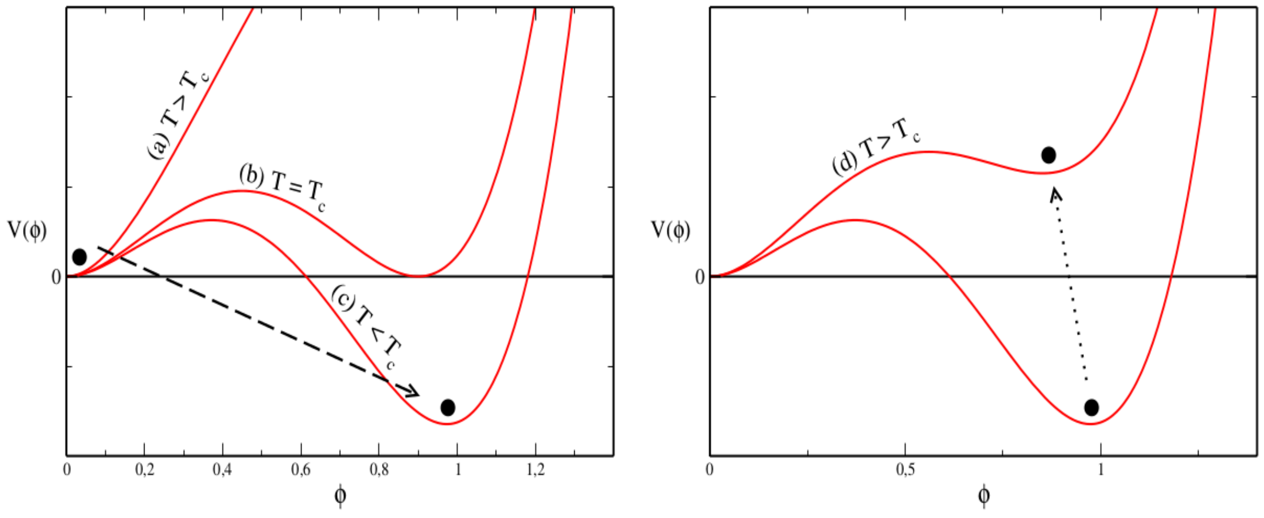


“Supersonic EWBG”

Thank you!

KEEP
CALM
AND
BACKUP
YOUR
WORK

Higgs Evolution in Early Universe → FINITE-TEMPERATURE EFFECTIVE POTENTIAL

$$\boxed{V_{\text{eff}}(h,T) = V_0(h) + V_0^{\text{loop}}(h) + V_T(h,T)}$$

Tree-level potential

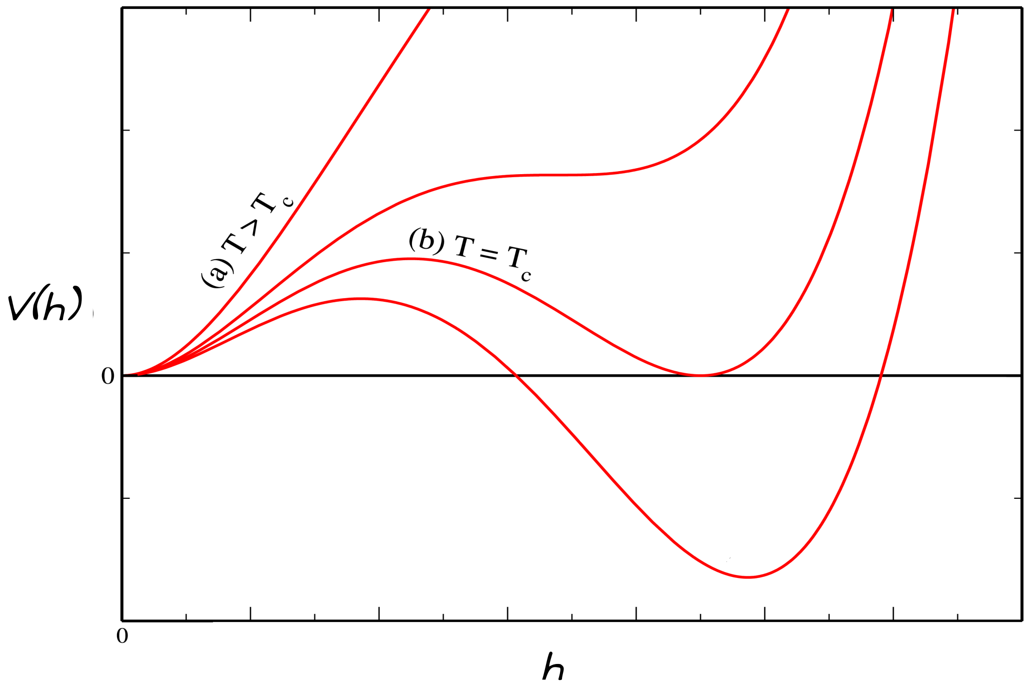
Loop corrections

Thermal corrections

(Perturbative) Nature of EWPT

1st Order:

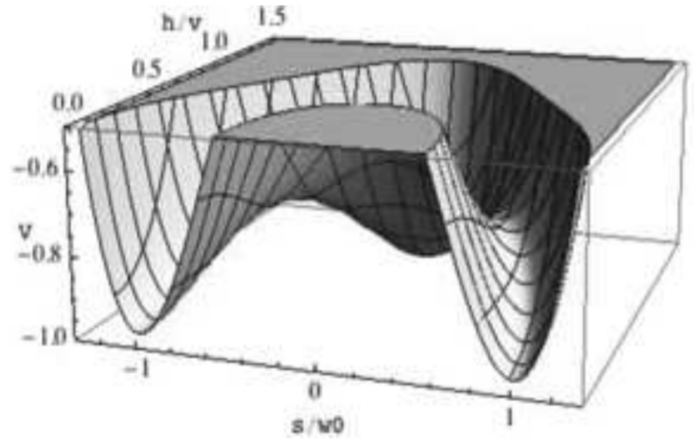
$\langle h \rangle = 0 \rightarrow \langle h \rangle = h(T)$ Discontinuous



Non-analytic term $(m^2)^{3/2}$ in $V(h,T)$ from Matsubara Zero-modes

(only present for bosons)

Multiple fields involved in the EWPT may allow for tree-level potential barrier

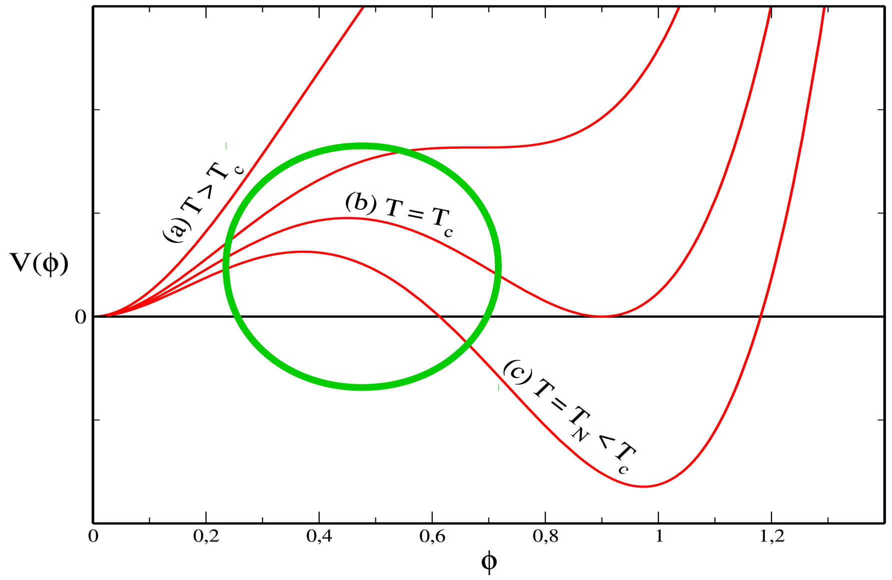


BSM: New Physics sizeably coupled to Higgs can drastically change the EWPT nature

- ▶ New Physics should induce deviations in Higgs couplings
- ▶ New Physics needed close to EW scale

Some further aspects of the EW Phase Transition

Effective Potential (finite T)



$$V_{\rm eff}(h,T) = V_0(h) + V_0^{\rm loop}(h) + V_T(h,T)$$

1-loop

$$V_1^T(h,T) = \frac{T^4}{2\pi^2}\left[\sum_i \pm n_i J_\pm\left(\frac{m_i^2(h)}{T^2}\right)\right]$$

$$J_\pm(x) = \int_0^\infty dy\, y^2 \log\left[1 \mp \exp\left(-\sqrt{x^2+y^2}\right)\right]$$

High-T expansion:

$$T^4 J_+\left(\frac{m^2}{T^2}\right) = -\frac{\pi^4 T^4}{45} + \frac{\pi^2 m^2 T^2}{12} - \boxed{\frac{T\pi (m^2)^{3/2}}{6}} - \frac{(m^4)}{32}\log\frac{m^2}{a_b T^2}$$

$$T^4 J_-\left(\frac{m^2}{T^2}\right) = \frac{7\pi^4 T^4}{360} - \frac{\pi^2 m^2 T^2}{24} - \frac{(m^4)}{32}\log\frac{m^2}{a_f T^2},$$

$$V_{\rm eff}(h,T) \approx (a\,T^2 - \mu^2)\,h^2 - E(T)\,h^3 + \lambda_{\rm eff}(T)\,h^4$$